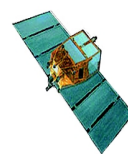




# The **BeppoSAX** Publications List





# The BeppoSAX X-Ray Astronomy Mission Publication List

## Introduction

The booklet includes a list of all the BeppoSax papers that have appeared in the scientific literature (refereed and non-refereed journals) updated to April 2016.

This list has been compiled using the Astrophysics Data System (ADS) searching for papers including either the word SAX, Beppo-SAX and BeppoSAX within the abstract. Although the list might be somewhat incomplete we believe that it includes the vast majority of the BeppoSAX papers. As the final contribution of this list we reproduce at the end of it the original paper on the scientific object of the mission (*Perola, G. C. et al. 1983*).

A comparison of the BeppoSAX scientific output with that of other satellites dedicated to space astronomy is shown in figure 1 where the rate of scientific publications (updated to August 2004) in refereed journals is plotted against time since launch.

Figure 1 shows that the maximum of the scientific output occurs a few years from launch. This maximum is then followed by a slow decay that depends on the satellite but can be roughly quantified with an halving time of 3-5 years even after the end of the mission.

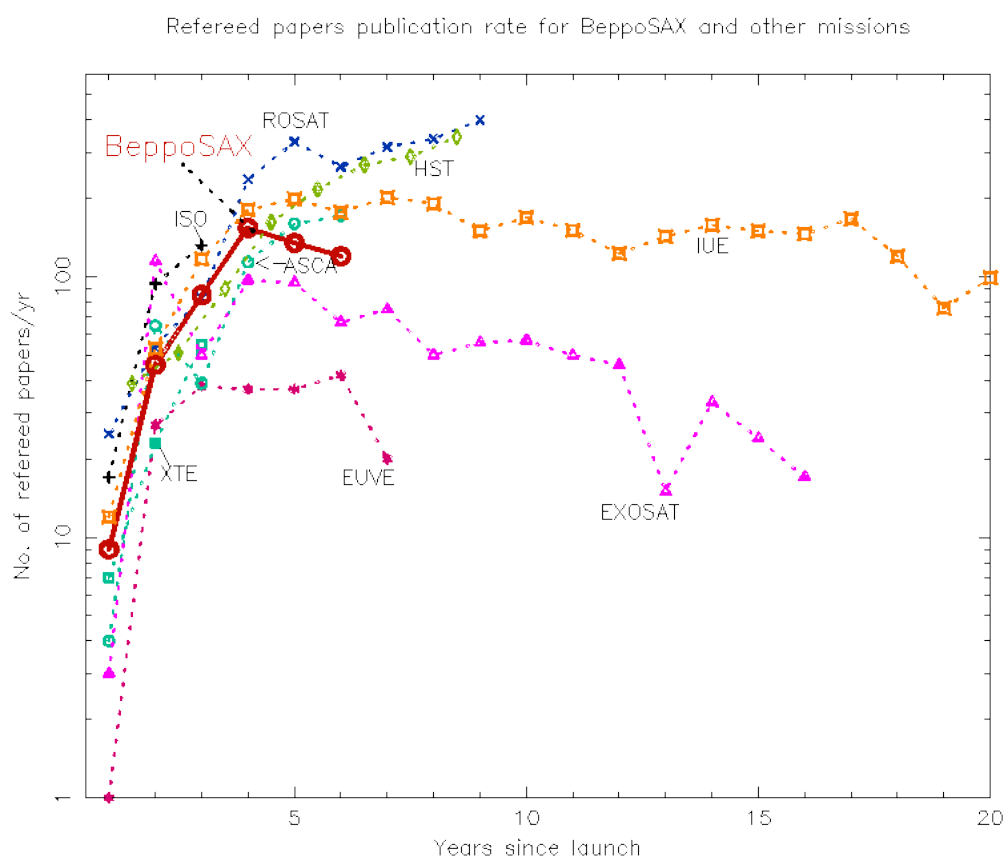


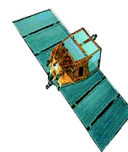
Fig. 1 The publication rate in refereed journals of several astronomy satellites is plotted against time since launch (update to August 2004)



## The BeppoSAX X-Ray Astronomy Mission Publication List

### Refereed publications

- [1] *Puccetti, S., A. Comastri, F. E. Bauer, W. N. Brandt, F. Fiore, F. A. Harrison, B. Luo, D. Stern, C. M. Urry, D. M. Alexander, A. Annuar, P. Arévalo, M. Baloković, S. E. Boggs, M. Brightman, F. E. Christensen, W. W. Craig, P. Gandhi, C. J. Hailey, M. J. Koss, S. La Massa, A. Marinucci, C. Ricci, D. J. Walton, L. Zappacosta, and W. Zhang; **Hard X-ray emission of the luminous infrared galaxy NGC 6240 as observed by NuSTAR**; *Astronomy and Astrophysics*; 585, A157 (2016)*
- [2] *Rivers, E., M. Brightman, S. Bianchi, G. Matt, K. Nandra, and Y. Ueda; **Suzaku confirms NGC 3660 is an unabsorbed Seyfert 2***; *Publications of the Astronomical Society of Japan*; (2016)
- [3] *Mineo, T., F. Massa, E. Massaro, and A. D'Ai; **Time properties of the the p-class burst of the microquasar GRS 1915+105 observed with BeppoSAX in April 1999***; *Astronomy and Astrophysics*; 586, A56 (2016)
- [4] *Vivekanand, M.; **The 1997 event in the Crab Pulsar in X-rays***; *Astronomy and Astrophysics*; 586, A53 (2016)
- [5] *Seifina, E., L. Titarchuk, and N. Shaposhnikov; **X-Ray Spectra of the High-mass X-Ray Binary 4U 1700-37 Using BeppoSAX, Suzaku, and RXTE Observations***; *The Astrophysical Journal*; 821, 23 (2016)
- [6] *Beri, A., B. Paul, M. Orlandini, and C. Maitra; **Temperature measurement during thermonuclear X-ray bursts with BeppoSAX***; *New Astronomy*; 45, 48 (2016)
- [7] *Marinucci, A., G. Matt, S. Bianchi, T. N. Lu, P. Arevalo, M. Baloković, D. Ballantyne, F. E. Bauer, S. E. Boggs, F. E. Christensen, W. W. Craig, P. Gandhi, C. J. Hailey, F. Harrison, S. Puccetti, E. Rivers, D. J. Walton, D. Stern, and W. Zhang; **The Seyfert 2 galaxy NGC 2110: hard X-ray emission observed by NuSTAR and variability of the iron K $\alpha$  line***; *Monthly Notices of the Royal Astronomical Society*; 447, 160 (2015)
- [8] *Kohn, S. A., M. J. Michałowski, N. Bourne, M. Baes, J. Fritz, A. Cooray, I. De Looze, G. De Zotti, H. Dannerbauer, L. Dunne, S. Dye, S. Eales, C. Furlanetto, J. Gonzalez-Nuevo, E. Ibar, R. J. Ivison, S. J. Maddox, D. Scott, D. J. B. Smith, M. W. L. Smith, M. Symeonidis, and E. Valiante; **Far-infrared observations of an unbiased sample of gamma-ray burst host galaxies***; *Monthly Notices of the Royal Astronomical Society*; 448, 1494 (2015)
- [9] *Seifina, E., L. Titarchuk, C. Shrader, and N. Shaposhnikov; **BeppoSAX and RXTE Spectral Study of the Low-mass X-Ray Binary 4U 1705-44: Spectral Hardening during the Banana Branch***; *The Astrophysical Journal*; 808, 142 (2015)



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- [11] *Wells, A.; Reflections on Swift from the early years; Journal of High Energy Astrophysics; 7, 12 (2015)*
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- [14] *Ducci, L., S. Mereghetti, D. Götz, and A. Santangelo; Ten years of INTEGRAL observations of the hard X-ray emission from SGR 1900+14; Astronomy and Astrophysics; 583, A113 (2015)*
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- [16] *van Putten, M. H. P. M., C. Guidorzi, and F. Frontera; Broadband Turbulent Spectra in Gamma-Ray Burst Light Curves; The Astrophysical Journal; 786, 146 (2014)*
- [17] *Seifina, E., L. Titarchuk, and N. Shaposhnikov; Black Hole Mass Determination in the X-Ray Binary 4U 1630-47: Scaling of Spectral and Variability Characteristics; The Astrophysical Journal; 789, 57 (2014)*
- [18] *Zhang, B.; Studies on High Energy Radiation Mechanisms and Gamma-Ray Burst Prompt Emissions; Acta Astronomica Sinica; 55, 354 (2014)*
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- [20] *Puccetti, S., A. Comastri, F. Fiore, P. Arévalo, G. Risaliti, F. E. Bauer, W. N. Brandt, D. Stern, F. A. Harrison, D. M. Alexander, S. E. Boggs, F. E. Christensen, W. W. Craig, P. Gandhi, C. J. Hailey, M. J. Koss, G. B. Lansbury, B. Luo, G. M. Madejski, G. Matt, D. J. Walton, and W. Zhang; The Variable Hard X-Ray Emission of NGC 4945 as Observed by NuSTAR; The Astrophysical Journal; 793, 26 (2014)*



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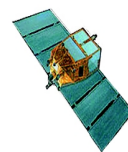
- [21] *Khabibullina, M. L., O. V. Verkhodanov, and V. V. Sokolov*; **Statistics of the Planck CMB Signal in direction of gamma-ray bursts from the BATSE and BeppoSAX catalogs**; *Astrophysical Bulletin*; 69, 472 (2014)
- [22] *Campana, R., M. Orlandini, E. Del Monte, M. Feroci, and F. Frontera*; **The radiation environment in a low earth orbit: the case of BeppoSAX**; *Experimental Astronomy*; 37, 599 (2014)
- [23] *Braito, V., L. Ballo, J. N. Reeves, G. Risaliti, A. Ptak, and T. J. Turner*; **Decoupling absorption and continuum variability in the Seyfert 2 NGC 4507**; *Monthly Notices of the Royal Astronomical Society*; 428, 2516 (2013)
- [24] *Piconcelli, E., G. Miniutti, P. Ranalli, C. Feruglio, F. Fiore, and R. Maiolino*; **Suzaku reveals X-ray continuum piercing the nuclear absorber in Markarian 231**; *Monthly Notices of the Royal Astronomical Society*; 428, 1185 (2013)
- [25] *Qin, Y., E.-W. Liang, Y.-F. Liang, S.-X. Yi, L. Lin, B.-B. Zhang, J. Zhang, H.-J. Lü, R.-J. Lu, L.-Z. Lü, and B. Zhang*; **A Comprehensive Analysis of Fermi Gamma-Ray Burst Data. III. Energy-dependent  $T_{90}$  Distributions of GBM GRBs and Instrumental Selection Effect on Duration Classification**; *The Astrophysical Journal*; 763, 15 (2013)
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- [30] *Chiang, C.-Y., E. M. Cackett, P. Gandhi, and A. C. Fabian*; **Confirmation of the nature of the absorber in IRAS 09104+4109**; *Monthly Notices of the Royal Astronomical Society*; 430, 2943 (2013)



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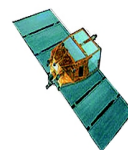
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- [32] *Kettula, K., J. Nevalainen, and E. D. Miller; EPIC using galaxy clusters*; Astronomy and Astrophysics; 552, A47 (2013)
- [33] *Bhattacharya, D., R. Misra, A. R. Rao, and P. Sreekumar; The uncorrelated long-term  $\gamma$ -ray and X-ray variability of blazars and its implications on disc-jet coupling*; Monthly Notices of the Royal Astronomical Society; 431, 1618 (2013)
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- [36] *Mück, B., S. Piraino, and A. Santangelo; A BeppoSAX and XMM-Newton spectral study of 4U 1735-44*; Astronomy and Astrophysics; 555, A17 (2013)
- [37] *Massa, F., E. Massaro, T. Mineo, A. D'Ai, M. Feroci, P. Casella, and T. Belloni; The complex behaviour of the microquasar GRS 1915+105 in the p class observed with BeppoSAX. III. The hard X-ray delay and limit cycle mapping*; Astronomy and Astrophysics; 556, A84 (2013)
- [38] *Banerjee, I. and B. Mukhopadhyay; Nucleosynthesis in the Outflows Associated with Accretion Disks of Type II Collapsars*; The Astrophysical Journal; 778, 8 (2013)
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- [42] *Mineo, T., E. Massaro, A. D'Ai, F. Massa, M. Feroci, G. Ventura, P. Casella, C. Ferrigno, and T. Belloni; The complex behaviour of the microquasar GRS 1915+105 in the p class*



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- [43] *Orlandini, M., F. Frontera, N. Masetti, V. Sguera, and L. Sidoli; BeppoSAX Observations of the X-Ray Pulsar MAXI J1409-619 in Low State: Discovery of Cyclotron Resonance Features;* *The Astrophysical Journal*; 748, 86 (2012)
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- [45] *Yaqoob, T.; The nature of the Compton-thick X-ray reprocessor in NGC 4945;* *Monthly Notices of the Royal Astronomical Society*; 423, 3360 (2012)
- [46] *Cackett, E. M., J. M. Miller, R. C. Reis, A. C. Fabian, and D. Barret; A Comparison of Broad Iron Emission Lines in Archival Data of Neutron Star Low-mass X-Ray Binaries;* *The Astrophysical Journal*; 755, 27 (2012)
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- [51] *Guidorzi, C., M. Lacapra, F. Frontera, E. Montanari, L. Amati, F. Calura, L. Nicastro, and M. Orlandini; GRBM;* *Astronomy and Astrophysics*; 526, A49 (2011)
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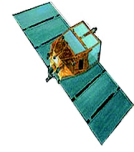


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- [55] *Cocchi, M., R. Farinelli, and A. Paizis*; **BeppoSAX view of the NS-LMXB GS 1826-238**; *Astronomy and Astrophysics*; 529, A155 (2011)
- [56] *Fiocchi, M., A. Bazzano, L. Natalucci, R. Landi, and P. Ubertini*; **The 17-min orbital period in the ultracompact X-ray binary 4U 0513-40**; *Monthly Notices of the Royal Astronomical Society*; 414, L41 (2011)
- [57] *Capitanio, F., A. J. Bird, M. Fiocchi, S. Scaringi, and P. Ubertini*; **Wide Field Camera: An In-depth Analysis**; *The Astrophysical Journal Supplement Series*; 195, 9 (2011)
- [58] *Sakamoto, T., S. D. Barthelmy, W. H. Baumgartner, J. R. Cummings, E. E. Fenimore, N. Gehrels, H. A. Krimm, C. B. Markwardt, D. M. Palmer, A. M. Parsons, G. Sato, M. Stamatikos, J. Tueller, T. N. Ukwatta, and B. Zhang*; **The Second Swift Burst Alert Telescope Gamma-Ray Burst Catalog**; *The Astrophysical Journal Supplement Series*; 195, 2 (2011)
- [59] *Ferrigno, C., M. Falanga, E. Bozzo, P. A. Becker, D. Klochkov, and A. Santangelo*; **4U 0115 + 63: phase lags and cyclotron resonant scattering**; *Astronomy and Astrophysics*; 532, A76 (2011)
- [60] *Vignali, C., E. Piconcelli, G. Lanzuisi, A. Feltre, C. Feruglio, R. Maiolino, F. Fiore, J. Fritz, V. La Parola, M. Mignoli, and F. Pozzi*; **On the nature of the absorber in IRAS 09104+4109: the X-ray and mid-infrared view**; *Monthly Notices of the Royal Astronomical Society*; 416, 2068 (2011)
- [61] *Kawaharada, M., K. Makishima, T. Kitaguchi, S. Okuyama, K. Nakazawa, and Y. Fukazawa*; **Suzaku Constraints on Soft and Hard Excess Emissions from Abell 2199**; *Publications of the Astronomical Society of Japan*; 62, 115 (2010)
- [62] *Finger, M. H., N. R. Ikhsanov, C. A. Wilson-Hodge, and S. K. Patel*; **Spin-Down of the Long-Period Accreting Pulsar 4U 2206+54**; *The Astrophysical Journal*; 709, 1249 (2010)
- [63] *Masetti, N., R. Landi, V. Sguera, F. Capitanio, L. Bassani, A. Bazzano, A. J. Bird, A. Malizia, and E. Palazzi*; **The peculiar high-mass X-ray binary 1ES 1210-646**; *Astronomy and Astrophysics*; 511, A48 (2010)

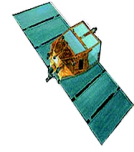




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- [65] *Horváth, I., Z. Bagoly, L. G. Balázs, A. de Ugarte Postigo, P. Veres, and A. Mészáros; Detailed Classification of Swift 's Gamma-ray Bursts; The Astrophysical Journal; 713, 552 (2010)*
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- [67] *Grange, Y. G., E. Costantini, J. de Plaa, J. J. M. in't Zand, F. Verbunt, J. S. Kaastra, and F. Verrecchia; X-ray spectral study of hot gas in three clusters of galaxies; Astronomy and Astrophysics; 513, A2 (2010)*
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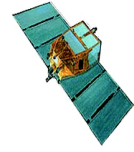
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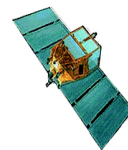
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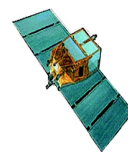
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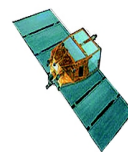
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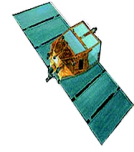


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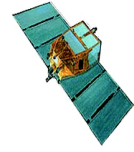
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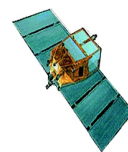
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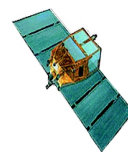
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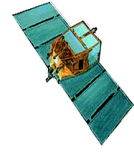
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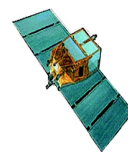
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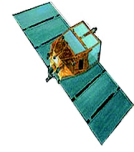


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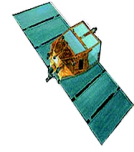




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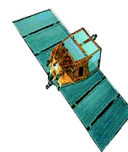
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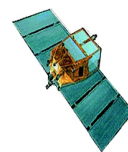
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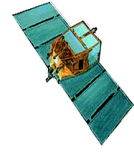
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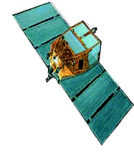
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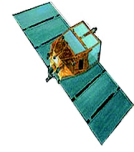


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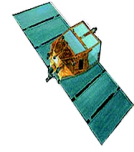




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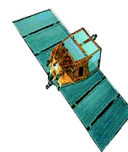
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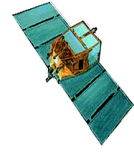
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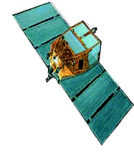
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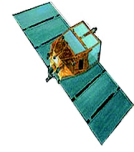


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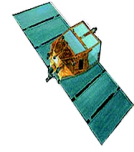




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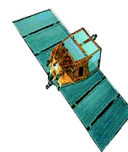
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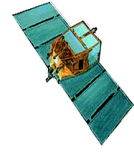
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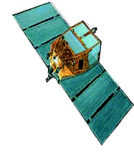
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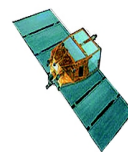


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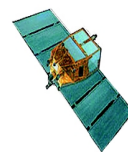




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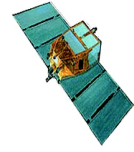
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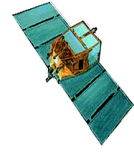
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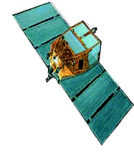
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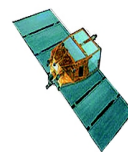


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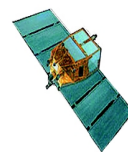




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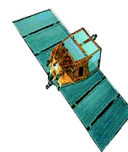
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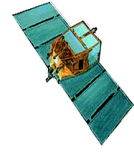
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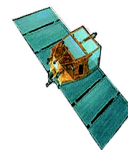
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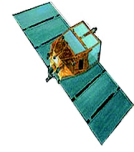
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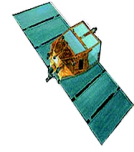




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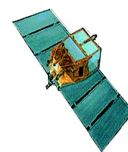
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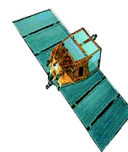
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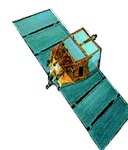
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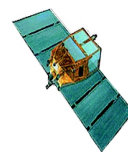
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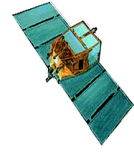
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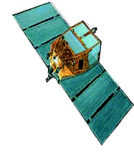
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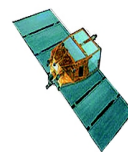


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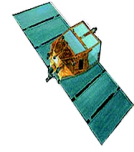




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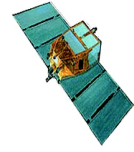
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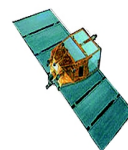
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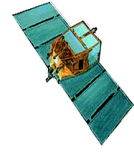
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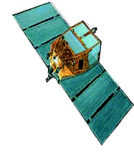
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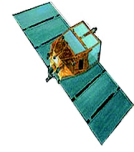


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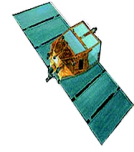




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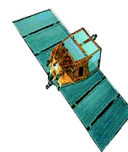
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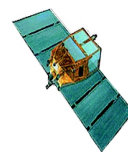
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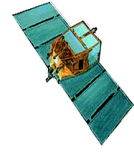
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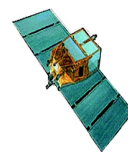
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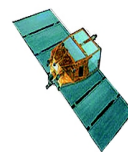


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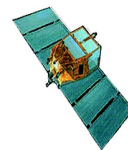
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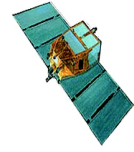
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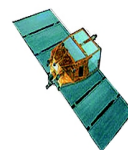
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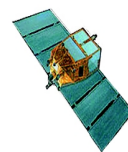
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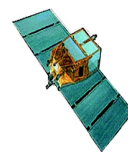
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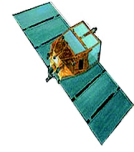


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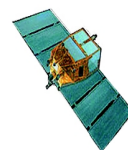
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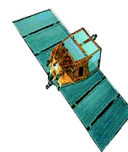
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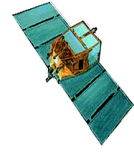


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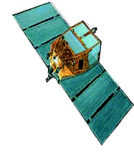
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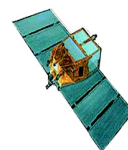


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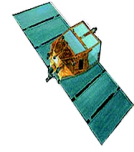




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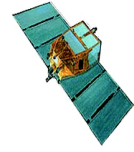
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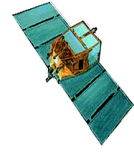
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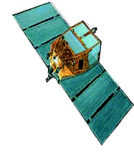
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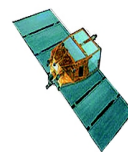


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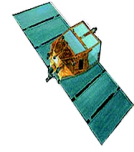




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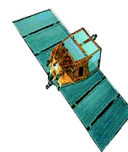
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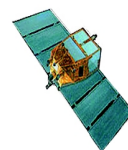
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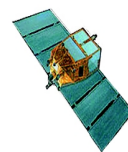
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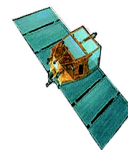
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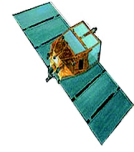
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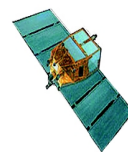




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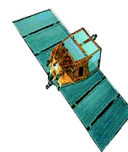
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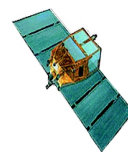
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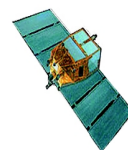
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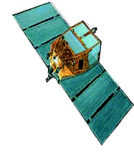
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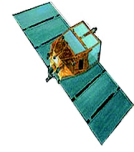


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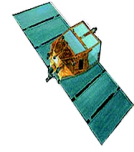




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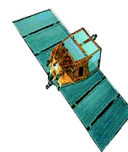
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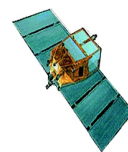
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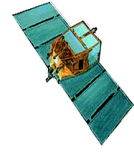
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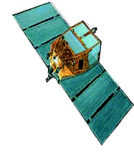
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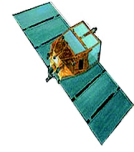


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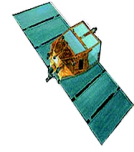




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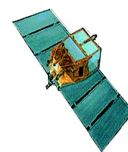
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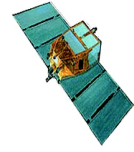
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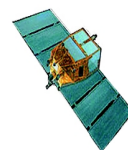
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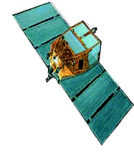
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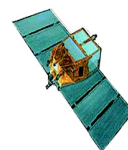


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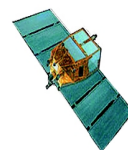




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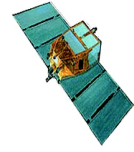
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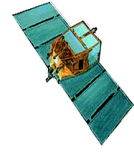
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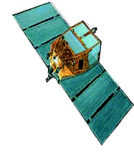
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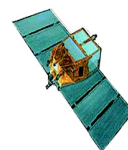


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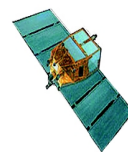




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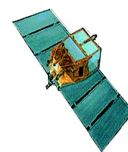
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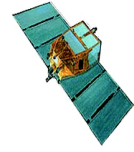
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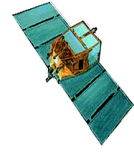
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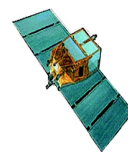
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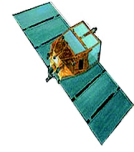


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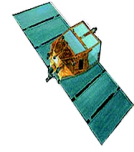




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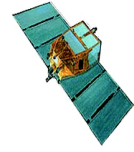
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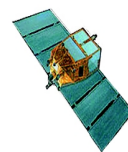
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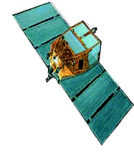
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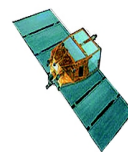


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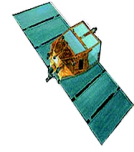




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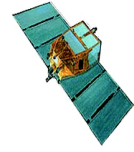
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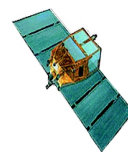
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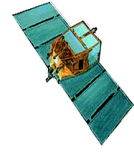
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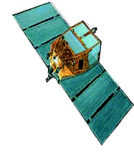
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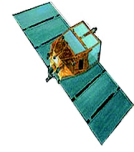


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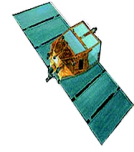




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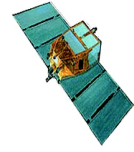
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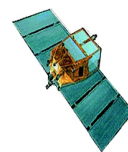
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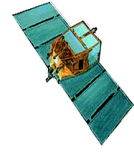
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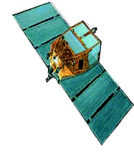
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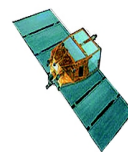


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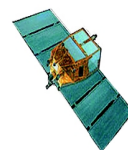




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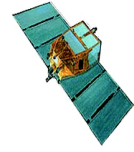
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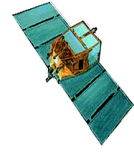
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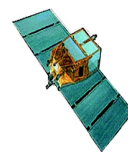
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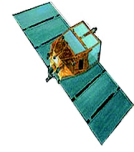


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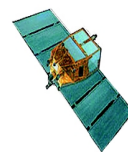




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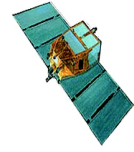
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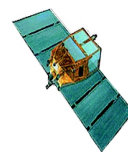
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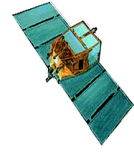
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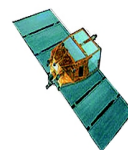
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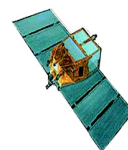


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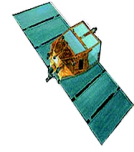




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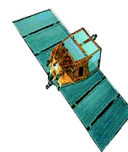
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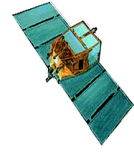
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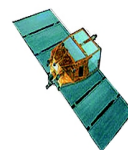
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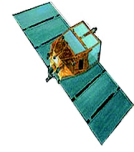


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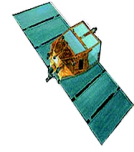




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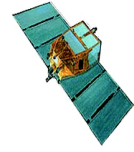
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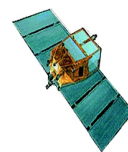
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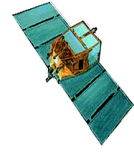
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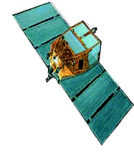
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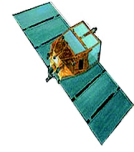


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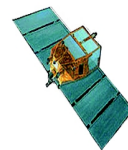
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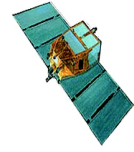
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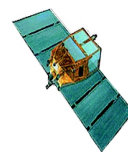
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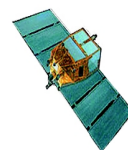
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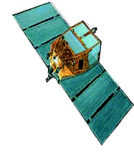
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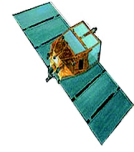


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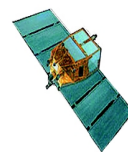




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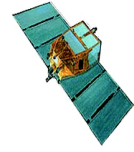
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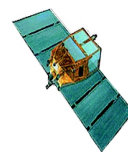
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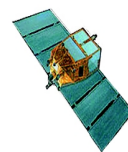
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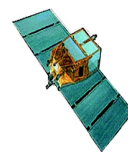
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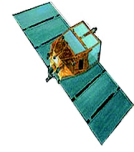


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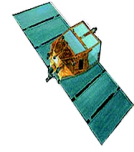




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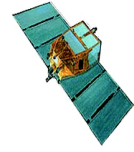
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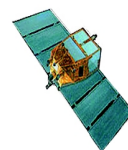
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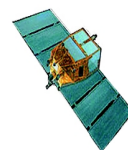
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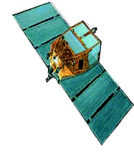
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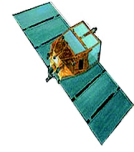


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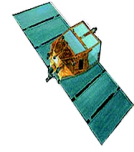




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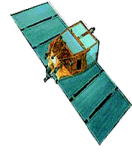
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# THE SCIENTIFIC OBJECTIVES OF THE SAX MISSION\*

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## 1. INTRODUCTION

High energy phenomena in astrophysics, involving either very high temperature plasmas or relativistic particles, are characterized by the emission of most or an important fraction of energy in the form of X rays. In one decade (1970-1981) a succession of X-ray astronomical satellites (Uhuru, ANS, SAS-3, Ariel V, OSO-8, Copernicus, HEAO-1 and the Einstein Observatory (EO)) has led to a large number of exciting and often unexpected results, and has consolidated the notion that this kind of phenomena are very common among the known galactic and extragalactic types of objects. The X-ray powers range from  $10^{27}$  erg/s of the Sun-like stars to  $10^{47}$  erg/s of the most extreme examples of activity in galactic nuclei, the quasars. Spectral, timing and imaging data have provided us with fundamental information on several topics, such as the active regions on the surface of the stars, the process of mass exchange and the nature and properties of the collapsed or degenerate objects in binary systems, the energy source and the structure of the emitting region in the active galactic nuclei, the interaction of supernova ejecta with the ISM, the distribution, chemical composition and mass of the hot intergalactic gas in clusters of galaxies, the origin of the diffuse X-ray background.

The missions up to HEAO-1 have been devoted to the observation mainly at medium X-ray energies (2-30 keV) of the brightest few hundred sources, spanning a range  $1:10^4$  in flux. The first satellite-borne X-ray telescope, the EQ, thanks to the imaging capability, has increased by a further factor  $10^3$  the flux range at the low X-ray energies (0.2-4 keV), and by more than one order of magnitude the number of known

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(\*) This presentation has been prepared with the cooperation of the following persons: A. Ferrari, F. Frontera, R. Fusco Fermiano, J. Heise, L. Maraschi, A. Preite Martínez, R. Robba, M. Salvati, A. Treves.

X-ray sources, thus performing for the first time an extensive study of classes of objects which are either intrinsically weak, or powerful but distributed over very large (cosmological) distances, and a detailed mapping of extended sources.

At high X-ray energies (50-200 keV) HEAO-1 has produced some remarkable results, but sufficiently deep and extensive investigations could not be carried out, mainly because of background problems and the need of very long exposure times. At these energies, balloon-borne experiments are still competitive, and have led to the important discovery of cyclotron lines in X-ray pulsars.

In addition to the broad band spectroscopy with the limited resolution typical of the commonly used proportional counters, with the EO a few higher resolution spectra could be obtained with the solid state detector (SSS), the grating and the focal plane crystal spectrometer, and only a rough spectral mapping of extended sources could be performed.

The two satellites now in operation, the Japanese TENMA and the European EXOSAT, both carry first generation gas scintillators performing broad band spectroscopy on bright sources with a resolution a factor of 2 better than proportional counters. With its low energy telescope and the large area proportional counters, EXOSAT is also performing broad band spectroscopy, timing and imaging mostly on a variety of known sources. The EO could only survey a tiny fraction of the sky ( $\leq 1\%$ ). The German X-ray telescope on ROSAT (scheduled for the mid 1987) will be devoted for the first part of the mission to a survey of the whole sky below 2 keV and for the second part to pointed and deeper exposures. The All Sky Survey alone is expected to detect a number of sources of the order of 100,000. The imaging with a grazing incidence telescope will be pushed to higher energies (10 keV) and better angular resolution ( $\leq 1$  arcsec) by the American AXAF, which is programmed for the 1990's.

Other missions due for the late eighties are the American XTE and the Japanese ASTRO-C, which will both emphasize fast timing with Large Area Proportional Counters. XTE will extend the energy range to 200 keV with the addition of a Phoswich. The Russian mission Saliut 7 (due in 1985-86) will also carry a composite payload, including a GSPC and a Phoswich.

Given this frame, the SAX mission was conceived to pursue scientific goals which appear strongly supported by the present knowledge and are to a large extent complementary to those of the missions scheduled for the late eighties.

## 2. SUMMARY OF SCIENTIFIC INSTRUMENTATION AND GOALS

The SAX payload consists of three instruments with a narrow field of view, and one instrument with a wide field of view. The three narrow

field instruments are:

- The Concentrator/Spectrometer (C/S), an assembly of four concentrators with position sensitive GSPC's in their focal planes, which combines mapping and broad band spectroscopic capabilities in the range, 0.1-10 keV, containing the most prominent atomic features typical of very hot ( $10^6-10^8$ °K) plasmas. With a spatial resolution of approximately 2' arc min, an effective area at 7 keV of  $\sim 175$  cm<sup>2</sup>, an energy resolution comparable to that of the SSS on the EO below 2 keV, and better by at least a factor of 2 than that of proportional counters at the iron K emission, this instrument will conduct detailed spectral studies of both point and extended sources down to very low flux levels. Its confusion limit is comparable to the limiting flux in the Deep Einstein Survey (Fig.1), the detection of a source at this limit requires about  $4 \times 10^4$  s of integration. The ability to investigate the iron K emission line and to detect short time variations is illustrated in Figg. 2 and 3.

- The Phoswich Detector System (PDS), an instrument for low resolution broad band spectroscopy in the range 15-200 keV. It will be used for detailed studies of the high energy continuum and its variations in the bright galactic and extragalactic sources, of the cyclotron features already discovered and for a search of features of this type in other binary pulsars. The timing and spectral capabilities of this instrument are illustrated in Fig. 4 and 5. The side CsI(Na) shields will be used as an all sky monitor of gamma-ray bursts, with a limiting sensitivity of  $10^{-6}$  erg/cm<sup>2</sup> s, and a timing capability from 0.5 to 10 ms, depending on the profile of the burst.

- The High Pressure Gas Scintillation Proportional Counter (HP GSPC), an instrument for moderately high resolution broad band spectroscopy in the range 3-120 keV. Its energy resolution is a factor of five better than that of the PDS, hence it will complement this instrument in the study of the cyclotron lines (Fig.5). In addition it will extend below 15 keV the spectral coverage of the PDS on the bright sources, and complement the C/S in the measurements of the iron K line (Fig.2).

A synthetic illustration of the sensitivity of the three coaligned narrow field instruments is given in Fig. 6.

The wide field instrument is made by:

- Three Wide Field Cameras (WFC), with a FOV of 27°x27° each, a spatial resolution of about 5' and an energy band 2-30 keV. They will be used for the long term monitoring of variable sources and detection/localization of transient sources, through a number of systematic surveys of the whole sky down to about 1 UFU as well as extended pointings on areas containing variable objects of particular interest, like the crowded galactic center region.

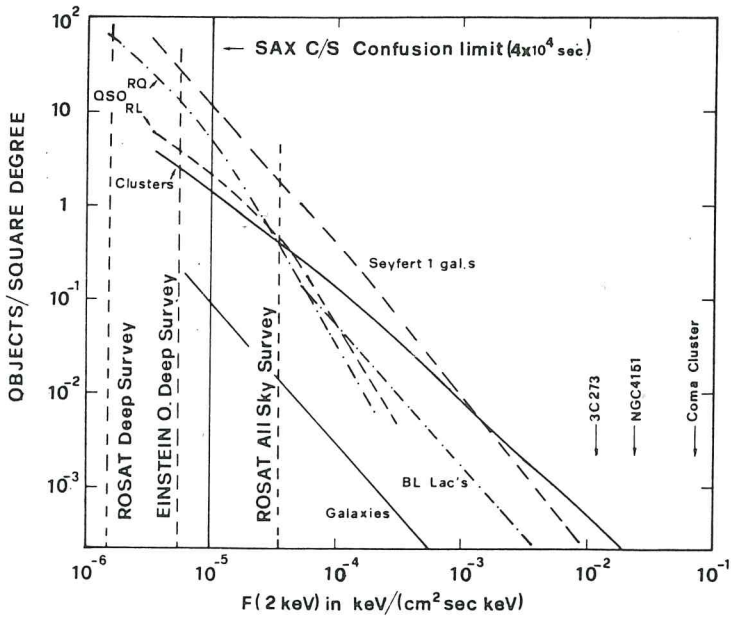


Fig. 1. Sensitivity-confusion limit of the C/S compared with Einstein and ROSAT. The predicted source counts are from Setti and Woltjer (1982).

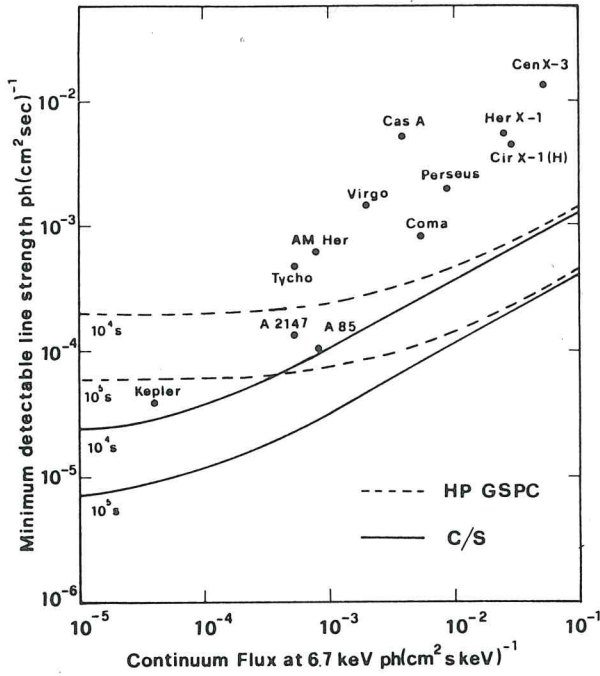


Fig. 2. Sensitivity of the C/S and HP GSPC to the 6.7 keV iron line.

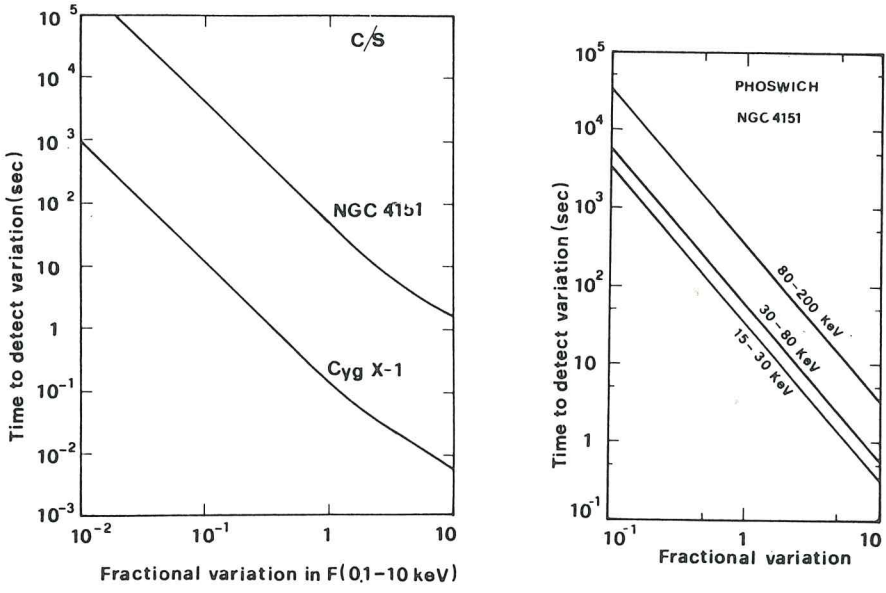


Fig. 3,4. Examples of the ability of the C/S and PDS to detect variations.

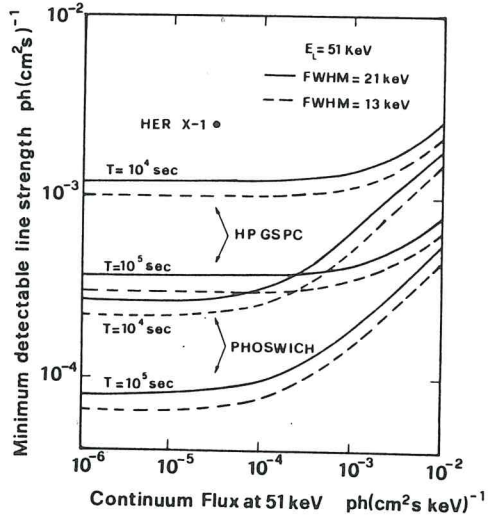


Fig. 5. Sensitivity of the PDS and the HP GSPC to a cyclotron line in emission with two different widths.



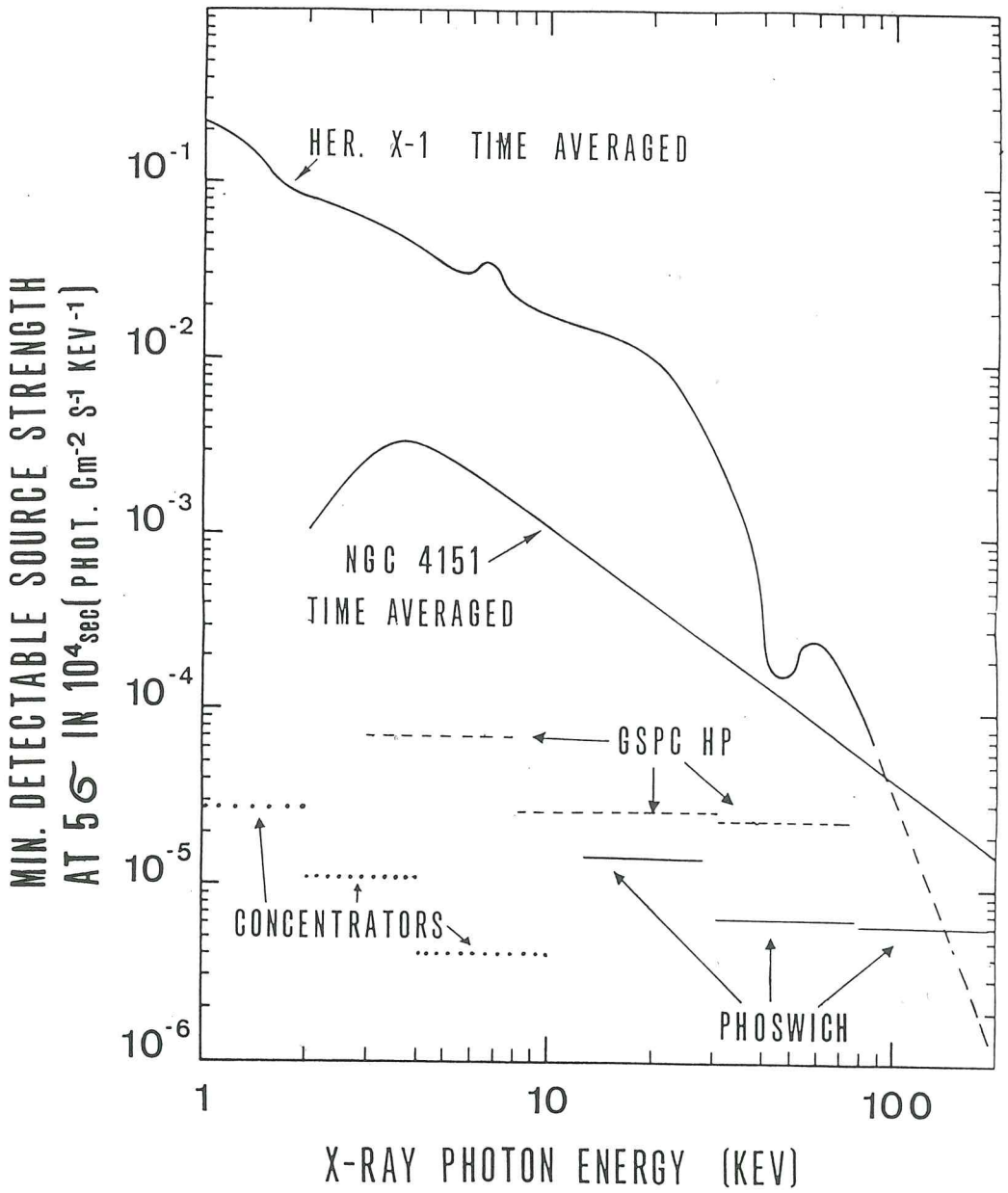


Fig. 6. Broad band sensitivity of the three Narrow Field instruments for an exposure time of  $10^4$  sec.

A detailed description of the four instruments and their performances is given in the accompanying paper II, by G. Spada.

### 3. SCIENTIFIC OBJECTIVES

In order to illustrate in more detail the effective capabilities of SAX, the present knowledge on various classes of X-ray emitting objects is summarized and the specific contributions expected from this mission are outlined.

#### 3.1. X-ray Binaries

Among the compact galactic X-ray sources, all those of high luminosity and a fraction of the low luminosity ones are thought to be binary systems powered by accretion of matter from a companion star onto a compact object, which may be a black hole, a neutron star or a degenerate dwarf. While the energy released is generally believed to be of gravitational origin (with the exception of the bursters), a variety of detailed theoretical models of the accretion process have been developed for the various types of objects belonging to this class, and further measurements are needed to improve our knowledge on this process. X-ray studies provide also information on the properties of the compact objects as well as on the evolution of binary systems.

A. X-ray binaries with an accreting black hole candidate. Cyg X-1 is the prototype of this class, its X-ray emission is characterized by a very conspicuous time variability over a wide range of time scales, from submillisecond to days, and by two distinct states, the so-called low states (the ordinary ones) and high states (the exceptional ones). The spectrum extends to several hundred keV, possibly to  $\sim 1$  MeV, with a smooth variable break in the 50-200 keV range and an excess tail above 500 keV (Fig.7). It becomes softer when the source undergoes a transition from LO to HI state. In addition, an ultrasoft component (if a black body,  $kT \sim 0.3$  keV) has been revealed, whose behaviour during transitions is not yet clear (Chiappetti et al, 1981). The spectral shape is generally interpreted, in the framework of accretion disk models, as the result of comptonization of soft photons by high temperature ( $kT = 50-80$  keV) gas (e.g. Shapiro et al, 1976). However the bimodal spectral behaviour is not well understood, and detailed studies during transitions of the spectral dynamics from 0.1 to at least 200 keV may provide the clue. These will be performed at epochs appropriately selected on the basis of a long term monitoring by the WFC.

A few more black hole candidates have been proposed, Cir X-1, GX339-4 (Fig.8) in our galaxy, LMC X-1 and LMC X-3 in the Large Magellanic Cloud (White and Marshall, 1983). Lately also the extraordinary object SS 433 has been added to the list (Grindlay et al, 1984). Their spectrum can

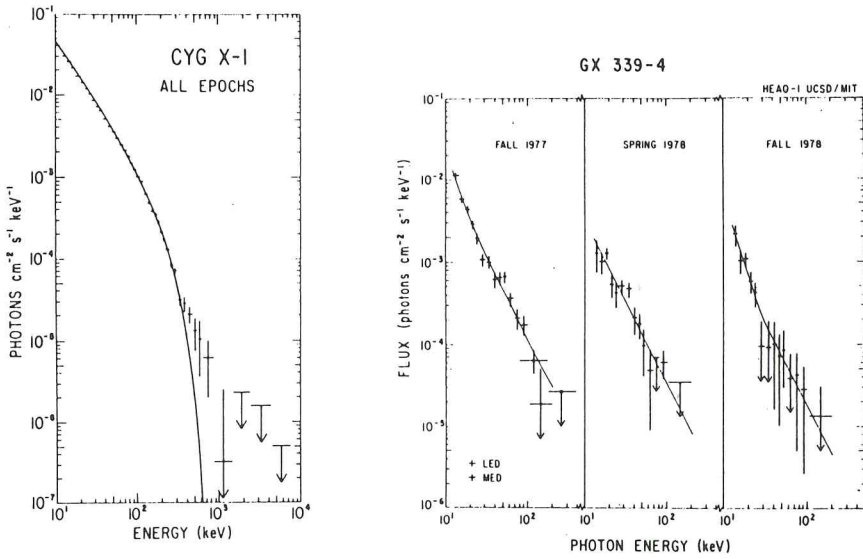


Fig. 7,8. Spectrum of Cyg X-1. The curve is a two-component comptonization model (Nolan et al, 1981). Spectrum of GX339-4 (Nolan, 1982).

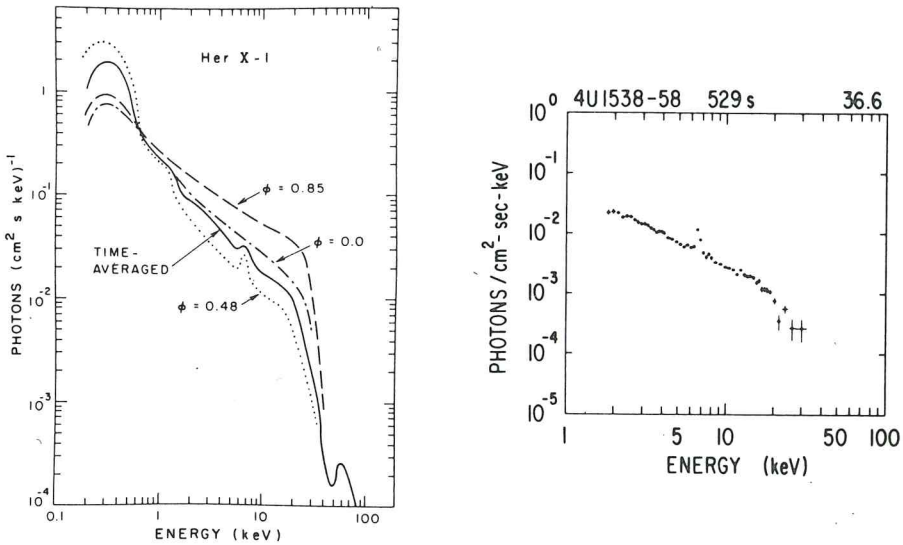


Fig. 9,10. Spectrum of Her X-1 at various pulse phases (McCray et al, 1982). Spectrum of a binary pulsar with a strong Fe line (White et al, 1983c).

be determined up to the 100 keV region in about  $10^4$  and  $10^5$  s for the galactic and LMC sources respectively. Furthermore, the iron line seen in Cir X-1 (Chiappetti and Bell Burnell, 1982), if confirmed, will be studied in detail (see Fig. 2).

The spectral measurements which will be performed on transient sources detected by the WFC may actually add more candidates, like the transient H1743-32, as suggested by White et al (1983a).

Comparative spectral studies in the wide energy range accessible to SAX may also provide a key to the understanding of the nature of the X-ray emission process in the AGN's (Sect. 3.6), as suggested, for instance, by White et al (1983b).

B. Accreting neutron stars with massive companions (X-ray pulsars). Since the discovery by Uhuru of the prototypes Cen X-3 and Her X-1, this class of objects has been the subject of the most extensive and detailed studies (e.g. White et al, 1983c, review by Joss and Rappaport, 1984). It comprises about 20 known X-ray pulsars representing  $\sim 15\%$  of the galactic sources with  $L_x \geq 10^{33}$  erg/s, some of which (about one quarter) are transients. The pulsation periods range from 69 ms to 835 s, and their derivatives imply a secular spin-up on time scales of  $10^2$ - $10^5$  years, but occasional spin-down's on comparatively short time scales have been observed. The orbital periods, measured in all but a few objects, range from a fraction of day to months. The X-ray spectra are generally hard power laws  $E^{-\alpha}$ , modified by the function  $\exp[(E_c - E)/E_f]$  at  $E > E_c$ , with typical values  $0.0 \leq \alpha \leq 0.5$  and  $10 \leq E_c \leq 20$  keV. In most cases the iron K line emission has been observed, with EW of several hundred eV. Features probably of cyclotron nature have been discovered in Her X-1 (Trumper et al, 1978), in the transient 4U0115+63 (Wheaton et al, 1979) and possibly in GX 1+4 (Maurer et al, 1982).

The picture emerging from the wealth of data collected so far (including the optical study of the companion and the determination of the orbital elements) is that of a strongly magnetized ( $B \geq 10^{12}$  G) rotating neutron star, which is accreting matter from the stellar wind of a massive companion and is secularly gaining angular momentum in the process. The sharpness of features in the pulse profiles is attributed to beaming in the emission, and is generally thought that both the accretion process and the emission mechanisms are strongly influenced by the magnetic field. The most relevant contributions expected from SAX on open problems are described in the following.

i) The continuum and its temporal behaviour. The X-ray pulsars display a variety of pulse shapes and continuum spectral variability across the pulse, especially at high luminosities, is often observed (Fig. 9,11). The pulse phase analysis as a function of energy yields fundamental information on the emission pattern, on the structure of

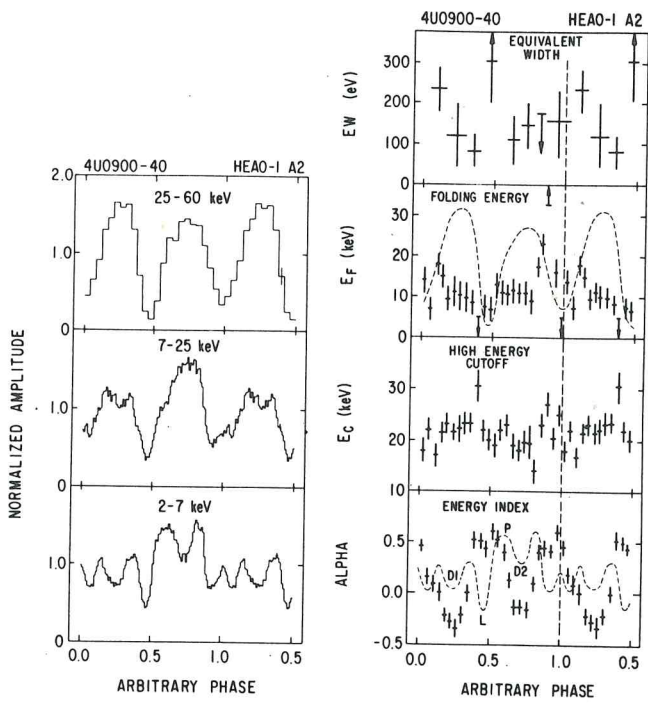


Fig. 11. The pulse light curve and variation of the spectral parameters with pulse phase of 4U0900-40. (White et al, 1983c).

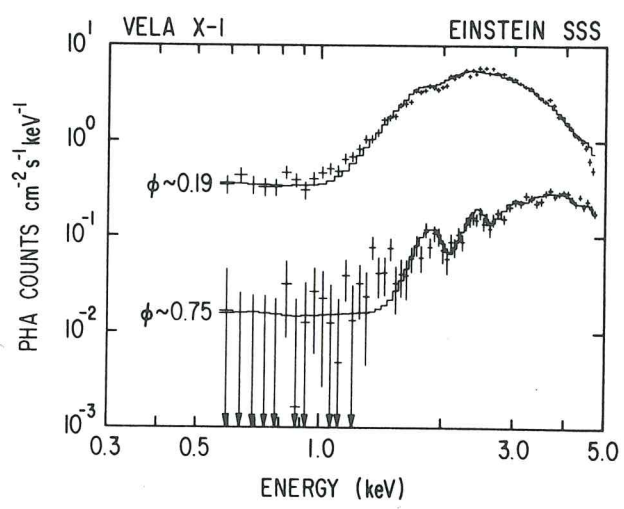


Fig. 12. SSS spectra of Vela X-1 at two values of the orbital phase (Kallman and White, 1982), showing the variation in the cutoff.

the accretion column, on the radiative transfer in the strongly magnetized plasma and on the dynamical interaction of the gas in accretion with the magnetic field and the outgoing radiation. The pulse shape and the spectral variability can be a function of the accretion rate as well as the orientation of the system to the line of sight, hence their dependence on the source intensity (comparison between low and high states in sources like Cen X-3, extended study of transient phenomena) and on the orbital phase are particularly informative.

In several cases (Fig.12) a photoelectric absorption by a column up to  $10^{23}$  H/cm<sup>2</sup> has been observed, which varies with the orbital phase, and is attributed to the stellar wind. In Her X-1 a soft X-ray component (Fig. 9) with pulse phase dependence is tentatively attributed to the reprocessing of hard X-rays in the beam intercepted by the inner edge of an accretion disk (McCray et al, 1982).

The detailed study of the spin-up trends and the spin-down episodes yields information on the angular momentum transfer mechanisms, and ultimately on the evolution of these systems. Furthermore the analysis of the pulse arrival times can give fundamental information on the internal structure of the neutron stars (e.g. Boynton, 1981).

The SAX mission can add significantly to our knowledge on these topics, through a combination of the long term monitoring with the WFC, aimed to detect transitions from low to high states, and sensitive observations with the pointing instruments of adequately selected fractions of the orbital period.

ii) Iron and other atomic features. In the majority of the known sources a conspicuous iron emission feature between 6.4 and 7 keV, with EW from 100 to 600 eV, has been observed (Fig.10, see HEAO-1 results in White et al, 1983c). In general only an upper limit of about 1.2 keV (FWHM) could be placed on the line width, but in some cases (e.g. Her X-1, Pravdo et al, 1977) evidence has been found of a broad or complex line structure. Variations in EW with the pulse phase and the luminosity state have been reported (Fig.11).

The feature is generally interpreted as fluorescence induced by the continuum either in the stellarwind or in the plasmasphere or both. The study of the fluorescence combined with that of the absorbing column and of absorption features (such as those at 2 and 2.5 keV detected in Vela X-1, Fig.12) as a function of the pulse and the orbital phases give information on the spatial distribution of the absorbing material and its ionization state, and should allow one to disentangle the stellar wind and the magnetospheric contributions.

With SAX, the factor 2-2.5 better resolution at 7 keV (compared with HEAO-1 data) will help in separating the iron fluorescence from

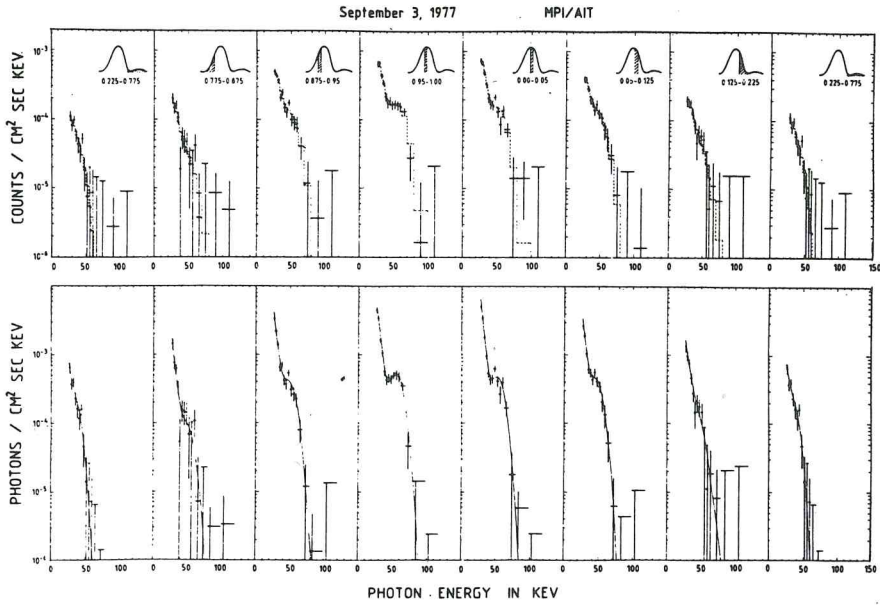


Fig. 13a. Phase dependent pulse minus off-source spectra of Her X-1. Lower panels: deconvoluted spectra with cyclotron line in emission.

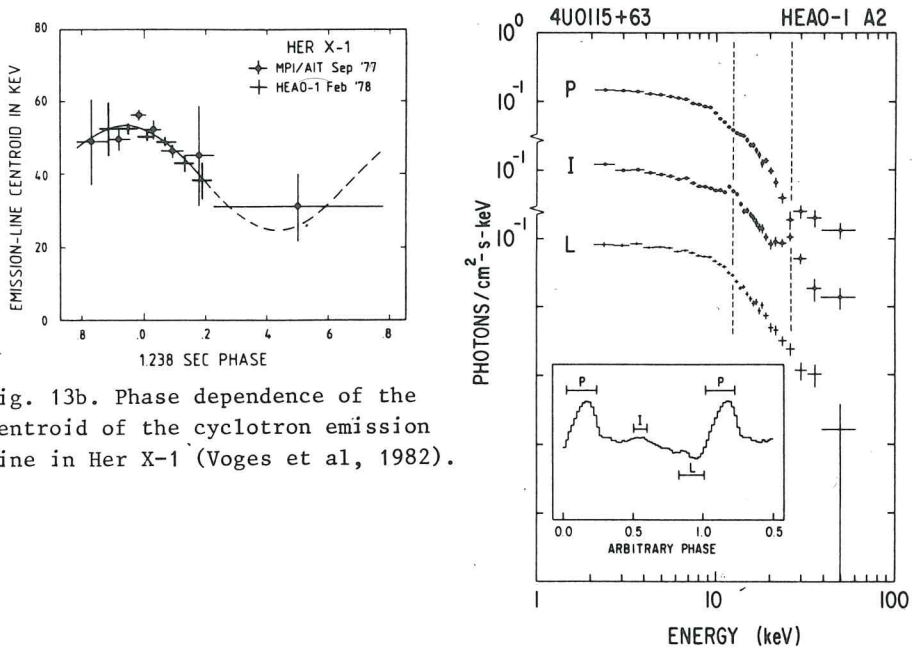


Fig. 13b. Phase dependence of the centroid of the cyclotron emission line in Her X-1 (Voges et al, 1982).

Fig. 14. Spectra of 4U0115+63 at three pulse phases, showing 1st and 2nd harmonic cyclotron features in emission/absorption (White et al, 1983c).

the corresponding absorption edge, and improve the information on the energy, the width and the structure of the line. Absorption and emission features below 3 keV will be studied with spectral resolution comparable to the one achieved with the SSS on the EO (Fig.12), but for much longer periods and more sources. From Fig. 2 one can appreciate the time needed for the detection of the iron line: in Cen X-3 and Her X-1 40 and 60 s are sufficient for detection, and a rough pulse phase analysis could be done already with an integration time of about 1000 s.

iii) Cyclotron lines. The study of emission/absorption features which can be identified as cyclotron lines is a powerful tool to measure the strength of the neutron star magnetic field and to probe the physical and geometrical parameters of the accretion process near the collapsed object through the radiation transfer effects. To this end one should be able to tell whether the line is in emission or in absorption and to measure its energy and width as a function of the pulse phase.

Present observational evidence indicate that: a cyclotron line may change with the phase from the emission to the absorption mode (Fig.14, 4U0115+63, where both the first and the second harmonics have been detected, White et al, 1983c); the feature may have a large intrinsic width and its centroid may vary periodically with the phase (Fig.13, Her X-1, Voges et al, 1982). The example of Her X-1 is also particularly instructive because the emission ( $E = 51$  keV) or absorption ( $E = 38$  keV) nature of the feature has not yet been observationally established beyond doubt, while theoretical models seem to favour the second alternative. The ambiguity in this situation is partly due to the limited photon statistics that could be obtained at high energies on the steeply declining continuum, and partly to the rather coarse energy resolution of the instruments.

To improve on the present knowledge, the pulse phase analysis must be applied to extended streams of high statistics low background data. An increase in energy resolution may be invaluable to avoid the ambiguity just mentioned, to investigate line profiles and to lower the detection limit for narrow features. The combination of the PDS and the HP GSPC then offers the opportunity to study in more details the cyclotron features discovered so far and to search for more of them in other objects of this class. Fig. 5 illustrates the detectability of a broad emission line at 51 keV, Fig. 15 compares the performance of the two instruments if the line is in absorption with a simulated observation of Her X-1.

C. Accreting neutron stars with low mass companions (including bursters). This class of objects, whose hystorical prototype is Sco X-1, is rather inhomogeneous and, with the 12 found in the core of globular



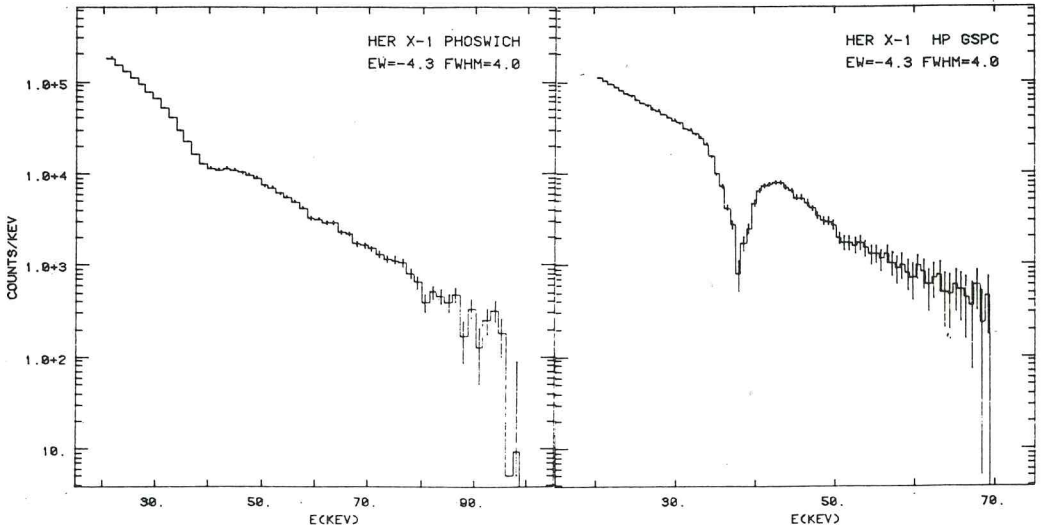


Fig. 15. Simulation of count spectra (on minus off-source) from the PDS and the HP GSPC obtained in  $10^5$ s exposure on Her X-1 in phase interval 0.0-0.1 around pulse peak, with cyclotron absorption on exponential continuum.

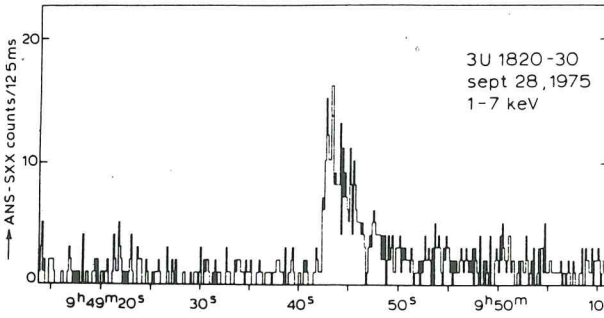
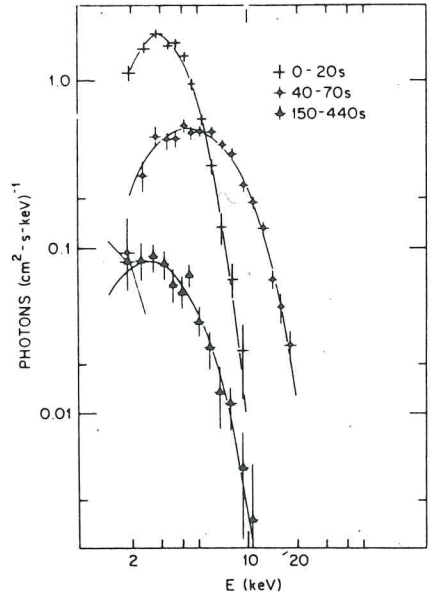


Fig. 16. One of the two discovery bursts in the globular cluster NGC 6624 (Grindlay et al, 1976).

Fig. 17. Spectra of a long burst taken with OSO-8 (Swank et al, 1977).



clusters, comprises about 50 known sources of high luminosity ( $L_x = 10^{35}-10^{38}$  erg/s), with a spatial distribution concentrated toward the galactic center (e.g. Bradt and McClintock, 1983). These X-ray sources do not pulsate (with the exception of 1627-673) and the evidence that the collapsed object is a neutron star is indirect. The mass transfer in these systems is believed to occur through Roche lobe overflow and to give rise to a large scale accretion disk, which often dominates the optical emission through the reprocessing of the X-rays.

The spectra are typically rather soft ( $kT \sim 4$  keV). Radiative transfer effects in an optically thick plasmasphere (such as the hot corona discussed by White and Holt, 1982) are thought to reshape the original source spectrum (and possibly to smear out the intrinsic pulsation). A strong iron line has been observed in Cyg X-3 (White and Holt, 1982), but is absent in all the other well exposed spectra.

The X-ray light curves show irregular variations along with, in some cases, periodic occultation phenomena, which can be used for the determination of the binary period (Walter et al, 1982) and for the study of the complex geometry of the matter in the source (White and Holt, 1982).

Very large variations, which have been seen in about half of the objects, take the shape of bursts of short duration (Fig.16), with typical rise times of about 1 second and decay times of several seconds to minutes (e.g. Lewin and Joss, 1983). In the majority of the cases the interval between bursts are of the order of hours to days (and possibly longer), and the spectrum softens during the decay, consistently with black body cooling (Fig.17). These bursts (called of type I) are attributed to thermonuclear flashes on the surface of the neutron star, as opposed to the conversion of gravitational energy responsible for the "steady" emission. In one object (4U1915-05) the bursts occur at intervals of seconds to minutes: these are called type II, and are thought to be powered by the same process as the "steady" emission.

An extensive and systematic monitoring of these sources will be possible with the WFC, through both periodic surveys and long exposures of the central region of the Galaxy, where the bursters in particular are concentrated. Coordinated optical observations could then be undertaken to observe the changes occurring as a consequence of the X-ray variations, in particular the optical bursts associated with the X-ray bursts (Pedersen et al, 1982). The pointing instruments, the C/S in particular, will study in detail the continuum spectra and their variation, search for and investigate the temporal behaviour of atomic features, like the iron line in Cyg X-3 and the oxygen edge discovered in Sco X-1 (Khan et al, 1981).

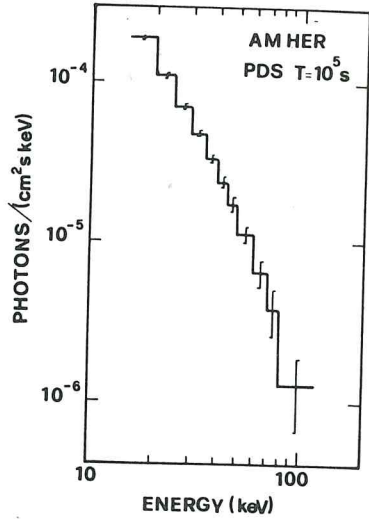
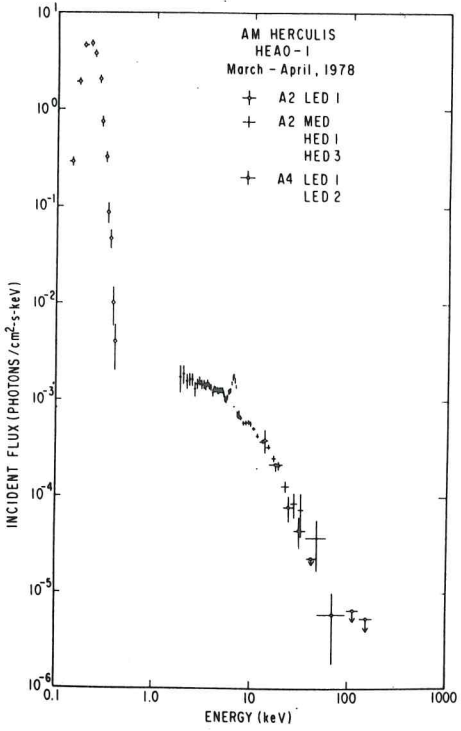


Fig. 18. Spectrum of AM Her, showing the soft and hard components and the strong iron emission line (Rotschild et al, 1981)

Fig. 19. The high energy spectrum of AM Her as could be obtained with the PDS in 10<sup>5</sup> s (if described by bremsstrahlung fit to the data in Fig. 18).

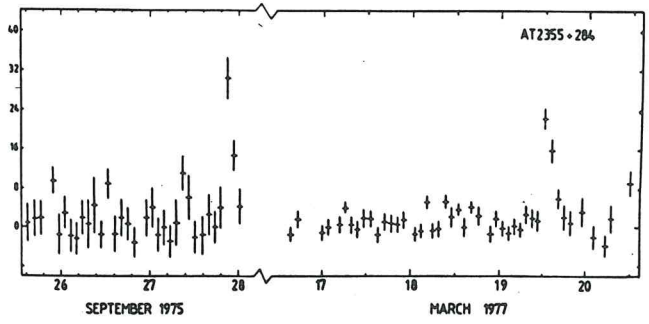


Fig. 20. The light curve of a fast transient measured with Ariel V (Pye and McHardy, 1983).

D. Accreting white dwarfs (cataclismic variables). Over 30 CV's are known to be X-ray emitters of relatively low luminosity ( $10^{29}$ - $10^{34}$  erg/s, Bradt and McClintock, 1983). These sources are believed to be close binary systems powered by accretion onto a white dwarf from a late-type star. The most powerful of them are the magnetic systems, like AM Her; where the matter is channeled by a strong field ( $B = 10^6$ - $10^8$  G) onto a very small portion of the white dwarf. The X-ray spectra (Fig.18) contain a very soft component in addition to a relatively hard one ( $kT \geq 5$ keV). The latter is interpreted as thin bremsstrahlung from the accretion shock, the former as black body emission from the heated stellar surface. The ratio of the power in the two components is very sensitive to the heat exchange processes and to the detailed structure of the accretion flow, and the present observational situation is rather intriguing (e.g. Heise et al, 1983).

The iron line has been observed with fairly large EW's, up to 800 eV (e.g. Rothschild et al, 1981). The light curves are characterized by intrinsic variations, such as recurrent outbursts lasting of the order of the week and pulsations with periods from seconds to minutes, and variations with the orbital phase due to eclipses.

Further insight on the physical properties of the accretion flow and the distribution of matter in these systems will be gained by means of very accurate determinations of the continuum, simultaneously at low and high energies, and of the iron line strength and structure, as a function of the high or low state of the output and of the orbital phase. Fig.19 illustrates the accuracy that can be achieved with a long exposure of the PDS on AM Her. The minimum time needed to detect the iron feature in the same source is about 300 s. The long term behaviour of the brightest sources will be studied with the WFC.

### 3.2. X-ray Transients (and Gamma Ray Bursts)

An X-ray source is classified as transient on the basis of its sudden appearance and limited duration above a minimum detectable flux. Their time scales range from a few minutes to several months, and most of them have been observed only once. A good fraction of the transient phenomena observed so far are now known to occur in conditions typical of well established classes of sources. The bright transients with long duration are typically confined to low galactic latitudes, and are recognized as binary systems with a highly variable rate of accretion onto a compact object. They have already been discussed in Sect. 3.1, and the WFC surveys will serve the purpose of detecting new and recurrent ones, which will then be studied more closely with the pointing instruments.

The weak, fast transients, which have been mainly discovered by

Ariel V (Pye and McHardy, 1983), are instead distributed isotropically over the sky (Fig.20-21). Most of them have lasted several hours, but events of a few minutes have also been observed. Among the 27 transients in the Ariel V sample the only major class of objects to be identified are six RS CVn-like systems, and 16 are not identified: they could be either nearby low luminosity galactic objects or extragalactic sources. The error boxes available prevent in fact the identifications to be made on the basis of a positional coincidence, notably most of the identifications proposed are with previously known X-ray sources. Moreover the spectral properties and the detailed temporal behaviour of this kind of transients could not be determined, due to the type of survey instrument used. One of the most specific tasks of the WFC will therefore be the detection, localization, temporal and spectral investigation of the high latitude transients, the number of detected events is expected to be about 100 in 2 years.

Another type of transient phenomena are the gamma-ray bursts. The first were observed more than 10 years ago, but their origin is still a mystery, partly because of the difficulty in obtaining accurate positions for a reliable optical identification. The PDS shield will be used as a gamma-ray burst monitor with a time resolution sufficiently good for a study of their history. About 20 bursts with energy  $> 10^{-7}$  erg/cm<sup>2</sup> are expected to fall in 2 years in the FOV of the WFC and be detected. Hence for a few of the bursts detected at gamma-ray energies a 5' position will be determined, in addition to the spectral and temporal study of the X-ray tail of their emission.

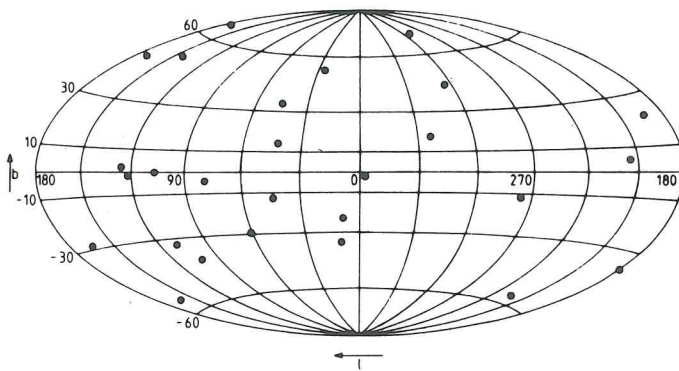


Fig. 21. Locations of the fast transient X-ray sources observed by the SSI instrument on Ariel V (from Pye and McHardy, 1983)

### 3.3. Supernova Remnants

The majority of the remnants of supernova explosions, like Cas A, are morphologically of shell type and their X-ray emission is of thermal origin. A minority of them, like the Crab Nebula, are center filled and emit radiation mostly of non-thermal (synchrotron) origin. There exist also remnants of intermediate type.

In the shell-type SNR's the energy radiated derives from the high velocity ejecta through the shock which develops in their interaction with the ISM. The X-ray spectra of some young (high temperature) remnants show strong atomic emission features: the  $K\alpha$ ,  $K\beta$  iron lines have been observed in particular with the HEAO-1, while the lines of highly ionized elements below 4 keV have been observed with the SSS on the EO (Fig. 22,23, e.g. Holt and McCray, 1982). The EO images of these remnants show a complex structure, indicative of propagation into an inhomogeneous ISM, but spatially resolved spectra are not yet available. The X-ray spectra of old remnants are more difficult to measure, because the surface brightness and the temperature are lower. A very informative spectrum in the 0.5 - 1.1 keV band, but only of a bright region of Puppis A, could be obtained with the FPCS on the EO (Winkler et al, 1981).

The interpretation of these spectra is complicated by the fact that, as predicted by theoretical models and confirmed by the observations, the temperature has a strong gradient across the shock structure and the ionization of the plasma is far from collisional equilibrium (e.g. Shull, 1982, Pravdo and Smith, 1979, Becker et al, 1980). Hence to derive, for instance, the chemical composition, and consequently fundamental information on the heavy elements enrichment of the ISM by SN explosions, it is mandatory to fit the spectral data with detailed model calculations, and, if the age  $t$  of the remnant is unknown, to obtain an estimate of  $t$  from the fit. The latter step is possible, at least for the remnants in the adiabatic phase, since the chemical composition itself has little effect on the evolution of the thermodynamical parameters.

To appreciate the value of the systematic studies, which could be carried out with SAX, it is convenient to follow the global approach outlined by Preite Martinez at this workshop. The hydrodynamical and ionization structure of a remnant are functions in general of the explosion energy  $E_0$ , the ISM density  $n$ , the age  $t$ . It can be shown that the four basic quantities  $E_0$ ,  $n$ ,  $t$ , and, if unknown, the distance  $d$ , can be related (analytically in the adiabatic phase) to four observable quantities, namely the continuum flux at an energy not contaminated by lines, the temperature  $T_c$  of the continuum derived from a bremsstrahlung fit at  $E \geq 4$  keV, the apparent size  $\theta$  and the centroid of the iron  $K\alpha$  blend. Other blends could in fact be used, but this one is the best candidate because it is well isolated and the iron itself is the only element

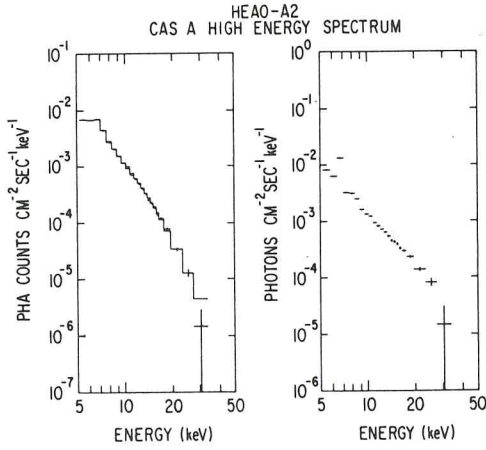


Fig. 22. Spectrum of Cas A showing the  $K\alpha$ ,  $K\beta$  Fe lines (Pravdo et al, 1979).

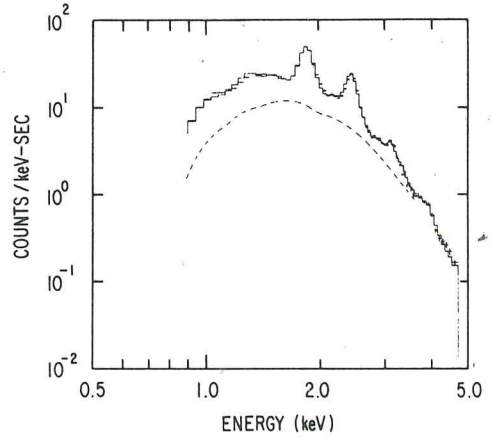


Fig. 23. Spectrum of Cas A obtained with the SSS (Becker et al, 1979).

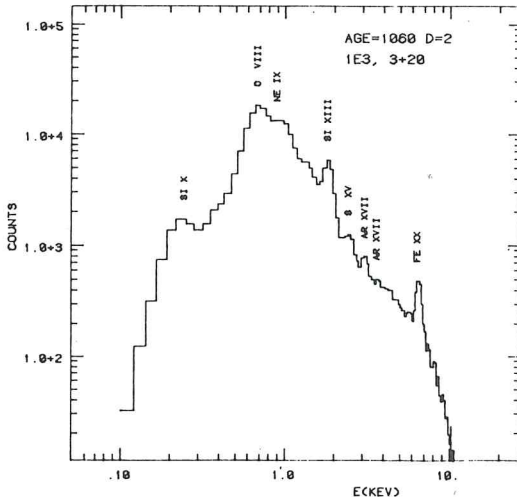
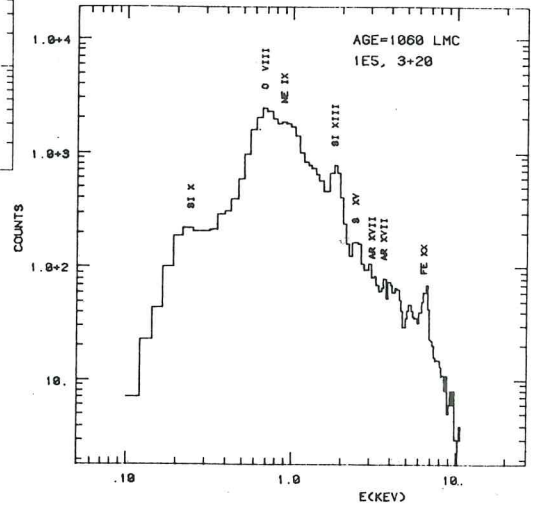


Fig. 24a. Simulation of count spectrum from the C/S in a  $10^3$  s exposure on model SNR of age 1060 y placed at 2 kpc.  $N_H = 3 \times 10^{20} \text{ cm}^{-2}$ .

Fig. 24b. Same as Fig. 24a, but with SNR placed at the distance of the LMC and  $10^5$  s exposure.



contributing to the blend. Once the basic quantities have been estimated in this way, the chemical abundances can be derived from the detailed model fit to the spectral data. Situations where either t (hystorical remnants) or d (e.g. the remnants in the Magellanic Clouds) or both are known, one or two of the observable quantities is redundant, and if all four can be measured accurately, the internal consistency of the model adopted can be tested. Discrepancies are expected between observations and predictions, as a consequence of inhomogeneities in the ISM and possible anisotropies in the explosion, and spatially resolved spectra should be obtained to evaluate their relevance.

The sensitivity to the iron line in SNR's can be estimated from Fig. 2. Due to the spectral complexity, more informative are simulations of the count/channel distributions for remnants with a typical value of  $E_0$  and n and different values of the age and the distance, based on a detailed model by Fusco Fermiano and Preite Martinez: examples are shown in Fig. 24-25, and they have been used to estimate the exposure times needed to carry out the following programs:

i) High statistic spectral study of the four hystorical remnants Cas A, Tycho, Kepler and AD 1006. Three of them have a radius  $> 2'$ , hence high quality spatially resolved spectra will be obtained.

ii) Spectral mapping of bright spots in nearby large SNR's, which are substantially older than the previous four. Some of them have probably past the adiabatic phase.

iii) Survey of a sample of radio SNR's. An important outcome of this program will be the determination of the distance, to be compared with the one derived from the  $\Sigma - D$  relationship.

iv) Detailed spectral study of a sample of SNR's in the Magellanic Clouds. About 30 remnants have been detected with the EO (Mathewson et al, 1983), and they represent an uniquely large sample of objects of different age at the same distance, for most of which statistically good spectra can be obtained with the C/S, with integration times  $\leq 10^5$  s.

The programs iii) and iv) include also some Crab-like remnants.

### 3.4. Stellar Coronae

One of the major discoveries of the EO was the detection of coronal emission from stars along the entire main sequence, of all luminosity classes, pre-main sequence stars as well as very evolved stars. The results from an extensive survey (Vaiana et al, 1981) contain over 140 detections with  $F(2\text{keV}) \geq 10^{-5}$  keV/(cm<sup>2</sup> s keV), while about 46% of the sources in the galactic plane above a limiting flux  $F(2\text{keV}) = 10^{-4}$  keV/(cm<sup>2</sup> s keV) are probably due to coronal emission from non-degenerate stars (Hertz and Grindlay, 1984). The X-ray luminosities are in the range  $10^{31}$ - $10^{34}$  erg/s for the early type OB stars,  $10^{26}$ - $10^{31}$  erg/s for the late type stars.



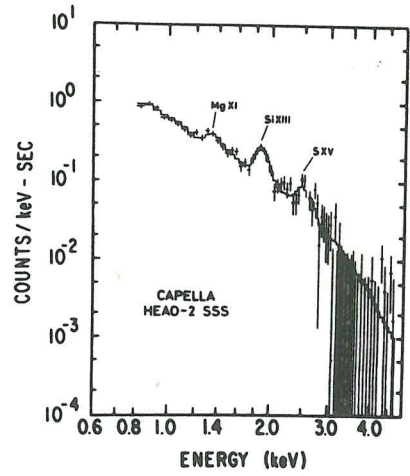
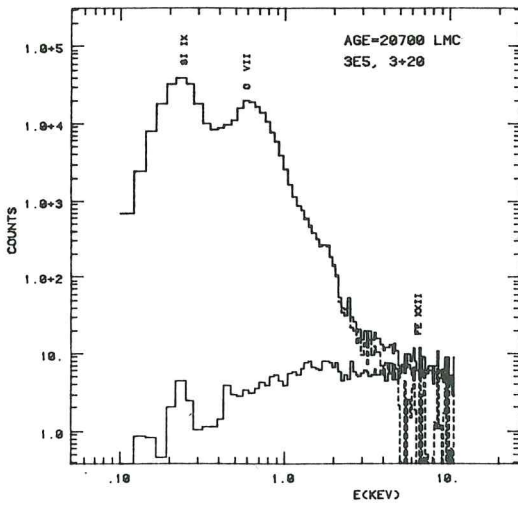


Fig. 25. Same as Fig.24, but with model SNR of age 20700 y at LMC distance, and exposure time  $3 \times 10^5$  sec. Lower counts are due to background.

Fig. 26. SSS spectrum of Capella (Holt et al, 1979).

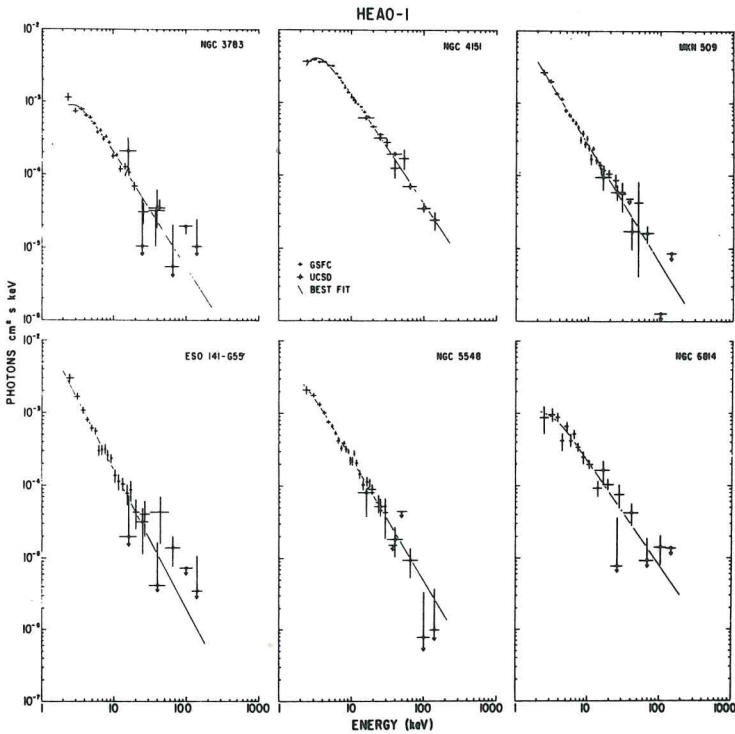


Fig. 27. High energy spectra of Seyfert galaxies (Rothschild et al, 1983).

The temperatures are typically in the range  $3 \times 10^6 - 10^7$  °K. The spectrum of Capella in Fig. 26 is a good example of the kind of information that can be obtained with the C/S on the temperature, density and chemical composition of the coronal plasma: in this particular case the fit required a blend of at least two temperatures and abundances of elements like Mg, Si, S and Fe different from the solar values (Holt et al, 1979). For the majority of the stars detected, however, even the estimate of the temperature is at best a rough approximation, on the other hand spectroscopic data are essential for the construction of detailed model atmospheres and the evaluation of the parameters which govern the coronal activity, a subject currently under very active investigation.

SAX can give an important contribution with the spectral study of a fairly large number of stars. In observations lasting  $2 \times 10^4$  s stars can be detected up to a distance  $D(\text{pc}) = 100(L_x/10^{29}\text{erg/s})^{1/2}$ , and spectra can be obtained, with a statistical quality comparable to that in Fig. 26 in the 0.8-2.5 keV range and better above 2.5 keV, of stars with  $F(2\text{keV}) \geq 10^{-3}$  keV/(cm<sup>2</sup> s keV). Above this limit there are about 10 stars in the survey by Vaiana et al (1981) and 50-100 are expected from a complete survey of the galactic plane ( $|b| < 15^\circ$ ), which should be detected by the All Sky Survey of ROSAT.

### 3.5. Normal Galaxies

The discrete X-ray sources in the Magellanic Clouds have been discussed together with the galactic sources in the previous sections. In the Andromeda Nebula 69 sources have been detected with an EO survey (van Speybroeck et al, 1979). They comprise isolated accreting binaries, globular clusters and SNR's, and have luminosities  $\geq 10^{37}$  erg/s. Notably the average luminosity of the (19) globular clusters in M31 is a factor of 7 higher than that of the (12) cluster sources in our Galaxy (Long and van Speybroeck, 1983), and the density of sources in the inner bulge of M31 is much greater than in our Galaxy. The spatial resolution of the C/S is insufficient to resolve the bulge sources, but for the other bright sources it will be possible to obtain spectral information for comparison with galactic sources of the same type, hopefully to find in this way some hints for the understanding of the remarkable difference in the bright tail of the luminosity function. Their detection requires exposures lasting 5000 s or less, hence exposures of the order of  $5 \times 10^4$  s should be sufficient to give useful spectral data.

The study of individual very luminous sources will be possible also in other members of the Local Group, whose distances are comparable to or less than that of M31 (0.7 Mpc).

For more distant galaxies only an investigation of their integrated emission will be feasible. Typical luminosities, as measured in a sample of 12 spiral galaxies with the EO (Fabbiano, 1983), range from  $10^{38}$  to

$10^{40}$  erg/s, and appear to correlate rather neatly with other (radio/optical) properties. At the confusion limit of the C/S the corresponding maximum distances range from 3 to 30 Mpc. Therefore it will be possible to extend the work done so far to measure the integrated spectra of a selection of normal galaxies out to at least the Virgo Cluster (20 Mpc). In addition to the correlations mentioned above, the spectral information may give further clues on the origin and the composition of their X-ray emission.

### 3.6. Active Galactic Nuclei

Objects of various types, Seyfert 1 and 2 galaxies (including radio galaxies with Seyfert-like optical spectra), QSO's (radio loud and radio quiet), BL Lac type objects, are classified as AGN's because they share the property of an intense non-stellar continuum, often extending from radio waves to hard X-rays, emitted from a very small region. The size of the region, inferred from the variability time scales, is generally a function of the energy band, and may be less than one light day. The X-ray emission is particularly important for the understanding of these phenomena at least for two reasons: the largest fraction of the power is often in the form of X-rays; the variability time scales in this band are the shortest among those measured so far.

The luminosity spreads over 6 orders of magnitude, from  $10^{41}$  erg/s of some nearby Seyfert 2 galaxies, to  $10^{47}$  erg/s of the most powerful QSO's. The X-ray spectra are typically hard power laws, with energy spectral index  $\alpha < 1.0$  and no obvious sign of a break in those measured up to 100-200 keV. Variability time scales as short as a few hundred seconds have been observed, but the most typical values are  $\geq 0.5$  day.

In the most popular theoretical models (e.g. Lightman, 1982) the power production occurs through the accretion of matter onto a very massive black hole ( $10^6$ - $10^8 M_{\odot}$ ), but the energy conversion process, the structure of the continuum source and the dominant radiation mechanisms are very controversial. By analogy with galactic sources containing black hole candidates, thermal emission with comptonization seems currently the favourite explanation of the X-ray spectra, but non-thermal mechanisms, like inverse Compton by relativistic electrons and even synchrotron radiation, cannot be excluded. Spectral and variability studies of the X-ray emission, coordinated with similar studies in the other observing windows, can be a powerful mean to unravel the structure of the source and the prevailing emission mechanisms.

From the optical spectra it is well known that the continuum source is surrounded by gas, which is responsible for the strong and broad emission lines. Photoelectric cutoffs and fluorescence emission features in the X-ray band, as observed in some objects, can be studied to probe the density and the spatial distribution of this gas.

The samples of known X-ray emitters in the three major categories are rather large: they contain over 40 Seyfert galaxies ( $10^{42}$ - $10^{45}$  erg/s, mainly from HEAO-1 and EO surveys, Piccinotti et al, 1982; Kriss et al, 1980), well over 100 QSO's, both radio loud and radio quiet, and out to a redshift  $z = 3.5$  ( $10^{43}$ - $10^{47}$  erg/s, mainly from EO surveys, Ku et al, 1980; Zamorani et al, 1981), and over 20 BL Lac objects ( $10^{43}$ - $10^{46}$  erg/s, mainly from EO observations, Schwartz and Ku, 1983). Notably a certain number of AGN have been discovered through their X-ray emission (e.g. Maccacaro et al, 1982). With these samples it has been possible to determine the luminosity function, to study the correlation between the X-ray and the optical/radio properties, to investigate the evolution with the cosmic time and finally to predict the number counts, as given in Fig. 1, for the different populations. On the basis of these statistical studies it has been shown that a large fraction at least of the X-ray background is due to the AGN's (e.g. Setti and Woltjer, 1982). Further insight into the XRB problem can be gained not only by increasing the statistics, but also by improving our knowledge on the X-ray spectra in the widest possible band for the various categories of AGN, as a function of power and redshift.

The C/S confusion limit, which corresponds to  $10^{-3}$  3C 273, is conveniently placed between the EO Deep Survey and the ROSAT All Sky Survey (Fig.1). Leaving to ROSAT the task of implementing the source count statistics, SAX will be mainly devoted to detailed spectral and variability studies of large representative samples of AGN's. The maximum distance to which the 0.1-10 keV spectra of objects with different powers are measurable can be judged from Table 1. The spectral studies up to 200 keV with the PDS will be necessarily restricted to relatively bright objects,  $F \geq 1/20(3C\ 273)$ .

Table 1. Maximum distance of AGN corresponding to detection limit of the C/S in  $3 \times 10^4$  s.

$L_x$ (0.5-4.5 keV)	Object	$z(\text{object})$	$z_{\text{max}} (q_0 = 0)$
$5.7 \times 10^{42}$	NGC 4151	0.003	0.1
$6.7 \times 10^{44}$	3C 390.3	0.057	0.9
$1.7 \times 10^{46}$	3C 273	0.158	3.2
$1.6 \times 10^{47}$	OH 471	3.400	5.9

A. Spectral studies of Seyfert galaxies. For a dozen Seyfert galaxies the spectra obtained with HEAO-1 instruments extend to 165 keV (Rothschild et al, 1983). Up to at least 50 keV the spectra are well represented by power laws with a very narrow distribution of spectral indices, mean value 0.67, r.m.s. deviation 0.15. No evidence of breaks above 50 keV has been found, but they cannot be ruled out, due to the limited statistics achieved (Fig. 27). For the bright NGC 4151 the measured spectrum extends to about 1 MeV, and there is some evidence of a variable break in the 100-200 keV region (Baity et al, 1983).

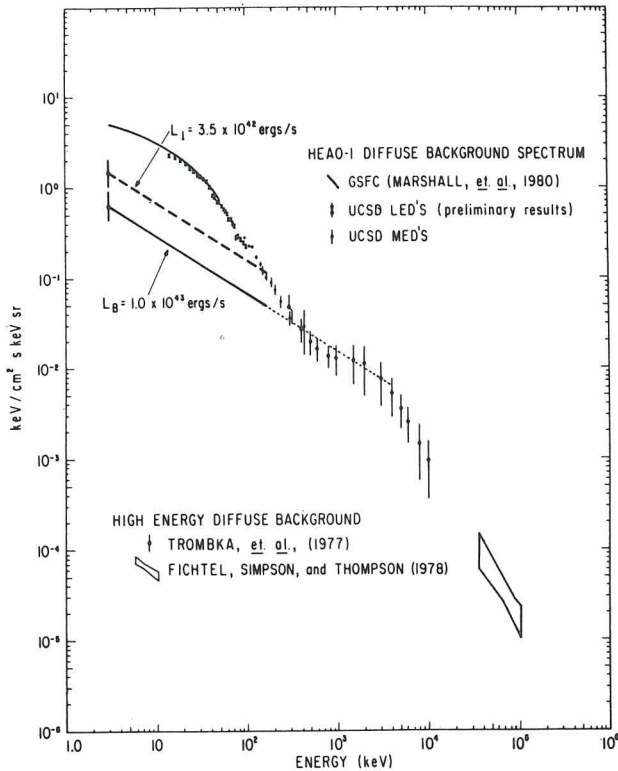


Fig. 28. The spectrum of the diffuse X-ray background. Shown is also the estimated contribution of the Seyfert galaxies for two choices of the form of the low luminosity break in their luminosity function. Their substantial contribution around 100 keV in the absence of spectral breaks is apparent (Rothschild et al, 1983).

With the PDS it will be possible to measure with good statistics the spectra of about 20 of the brightest Seyfert's up to the 80-120 keV band, and those of about 10 up to the 120-200 keV band, with an integration time  $< 3 \times 10^5$  s for each source. The search for a break, and its possible correlation with  $L_x$ , is important not only for the understanding of the emission mechanism, but also for a correct estimate of the contribution of this class to the XRB: according to Rothschild et al (1983) their estimated contribution may actually exceed the background above 100 keV, unless their average spectrum breaks in the region 50-150 keV (Fig. 28). Spectra of fainter Seyfert's will be measured with the C/S, to study the correlation between  $L_x$  and the slope and possible evolutionary trends with cosmic time.

Strong intrinsic photoelectric cutoffs are more common among the Seyfert's of low luminosity, and in some cases evidence of variability has been found (e.g. Mushotzky et al, 1980). In the objects NGC 4151, Cen A and NGC 5548 an iron emission feature has been detected, which, at least in the first two cases, has been interpreted as fluorescence. Combining spectral measurements of NGC 4151 with the EO-SSS and the HEAO-1-A2 instruments (Fig. 29), Holt et al (1980) were able to draw important conclusions on the spatial distribution of the cold gas around the source. The fluorescence has been detected so far together with a strong cutoff, but it should be present also if the line of sight happens to be clear of absorbing gas. The C/S spectral studies of a considerable sample of objects over a wide range of  $L_x$  should therefore add substantially to our knowledge on the density and geometry of the cold gas in AGN. Furthermore, as suggested for the apparently variable iron feature in NGC 5548 by Fabian and Ross (1981), they may also provide information on the hot gas component. An iron line with the EW = 200 eV measured in NGC 4151 could be detected in  $10^7$  s out to  $z = 0.25$  in a Seyfert 100 times more luminous.

B. Spectral studies of Quasi Stellar Objects. Only a few spectra have been measured between 2 and 40 keV, the best measurement (3C 273) is fitted by a power law with spectral index 0.4 (Fig. 30, Worrall et al, 1979). For the large sample detected with the EO the estimates from the hardness ratio are consistent with an average slope 0.4-0.5 (Zamorani et al, 1981). Since these objects, and the radio quiet in particular, may produce the major fraction of the AGN contribution to the XRB, statistical considerations imply that their average spectrum should be close to the slope of the XRB (0.4) below 40 keV, and break between 40 and 100 keV. Of course the average spectrum may conceal differences as a function of parameters like  $L_x$ , the radio power, the optical luminosity and the redshift, hence detailed spectral studies of appropriately chosen samples are needed.

The PDS will be used to measure the spectra up to 100-200 keV of

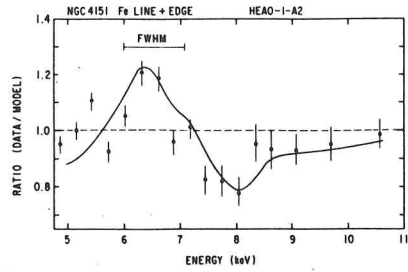
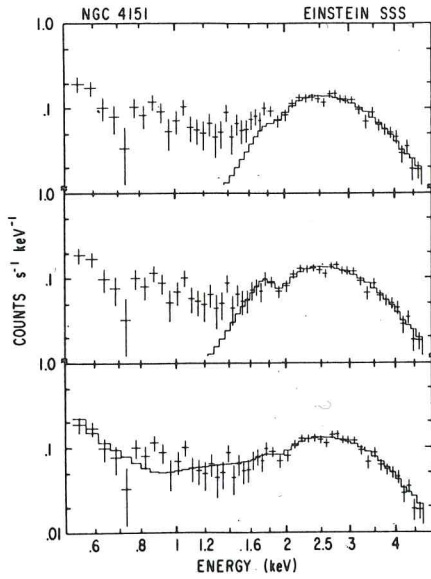


Fig. 29. NGC 4151. Left: low energy count rate spectrum with three trial fits. Up: evidence of iron edge and fluorescence (Holt et al, 1980).

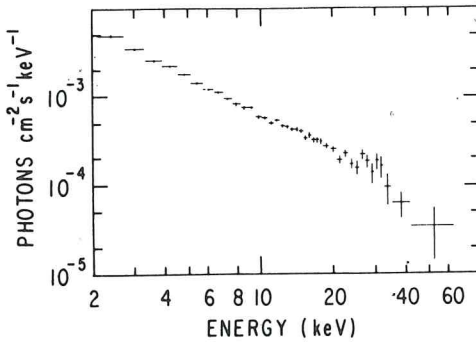


Fig. 30. HEAO-1-A2 spectrum of the quasar 3C 273 (Worrall et al, 1979).

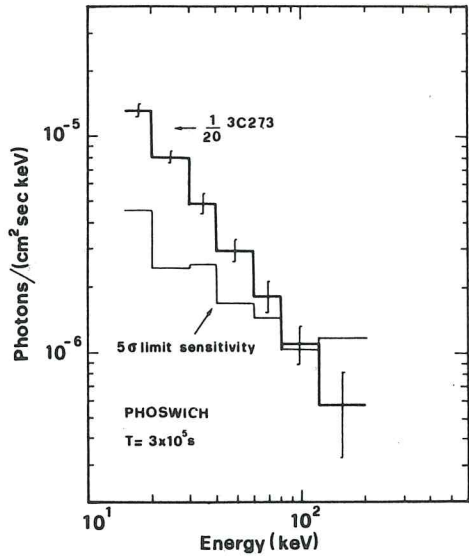


Fig. 31. Example of high energy spectrum obtainable with  $3 \times 10^5$  sec exposure of the PDS on QSO with same slope as 3C 273 but intensity 20 times lower.

about 10 bright QSO's, with typical exposure times  $10^5 - 3 \times 10^5$  s (Fig. 31), which will tell us about the presence of breaks above 40 keV at least in a small,  $z < 1$ , sample. The C/S can measure the spectral slope of a QSO with the luminosity of 3C 273 and with  $z = 3$  (hence up to 30-40 keV in its rest frame) with an exposure time of about  $5 \times 10^4$  s. To investigate the distribution of the spectral index as a function of the redshift and the other observable parameters, many QSO spectra could then be obtained in a reasonably short time. This information will be of great relevance also for the understanding of the emission mechanisms. For instance, the tight correlation between  $L_x$  and the power in the mm waveband, which holds for the radio loud QSO's (Owen et al, 1981) may be regarded as evidence that their X-ray emission contains a substantial inverse Compton contribution. On the other hand in some quasars (as well as BL Lac's) the X-ray flux appears to be significantly smaller than theoretically predicted from the measured radio parameters, a discrepancy which is usually attributed to relativistic beaming of the non-thermal source (e.g. Schwartz and Ku, 1983). With the SAX spectral measurements a further test of the IC model will be possible, through a comparison of the observed with the predicted slopes.

C. Spectral studies of BL Lac type objects. These objects differ from the other AGN's for the weakness of the optical emission lines, but share the properties of a high optical polarization and of a violent variability with the OVV quasars, and together they are often called "blazars". The few measured spectra show a large variance in the slope: a probable cause of this fact, illustrated in Fig. 32, is the existence of two spectral components, a hard one similar in slope to the other AGN's, and a soft one which could be the extrapolation of the synchrotron optical-ultraviolet continuum. It is generally believed that their continuum emission is enhanced in the observer frame by relativistic beaming, and the ratio  $L_x/L_r$  has been used to infer values of the bulk Lorentz factor up to 10-20.

Accurate spectra of many objects of this class could be measured with the C/S in a limited amount of time. For instance, less than  $10^5$  s in total are required for the 20 objects ( $z < 2.8$ ) listed by Schwartz and Ku (1983). The high energy spectra could be measured with reasonable accuracy only for less than 10 objects, but they will be most useful to disentangle the two components in the time devoted to the study of the variability.

D. Variability studies. The X-ray emission of the AGN's is subject to irregular intensity variations, occurring on time scales from a few hundred seconds (Fig. 33) to days (Fig. 34), months, years. The largest amplitude variations have been observed in blazars. The variability may be characterized by a frequent flare-like activity, isolated outbursts,



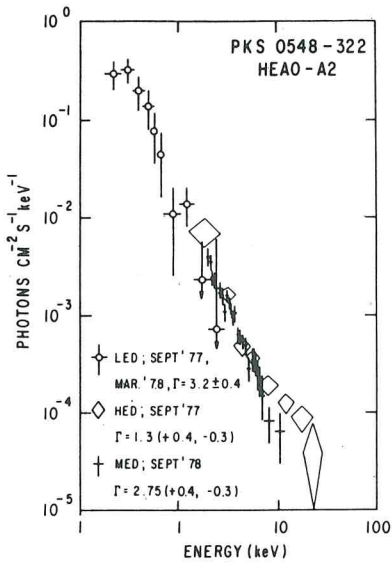


Fig. 32. Spectral variability in a BL Lac type object. The HED data points show the presence of a soft and a hard component (Worrall et al, 1981).

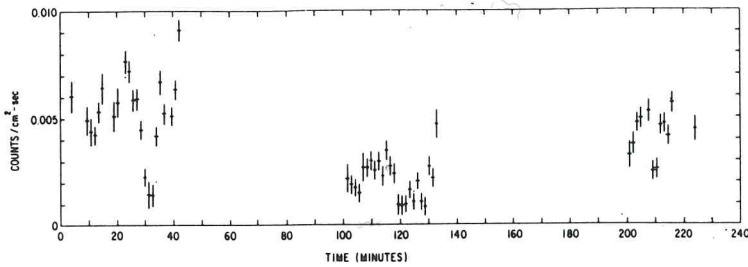


Fig. 33. Light curve of NGC 6814 from HEAO-1-A2 data, showing variability on the time scale of about 100 sec (Tennant et al, 1981)

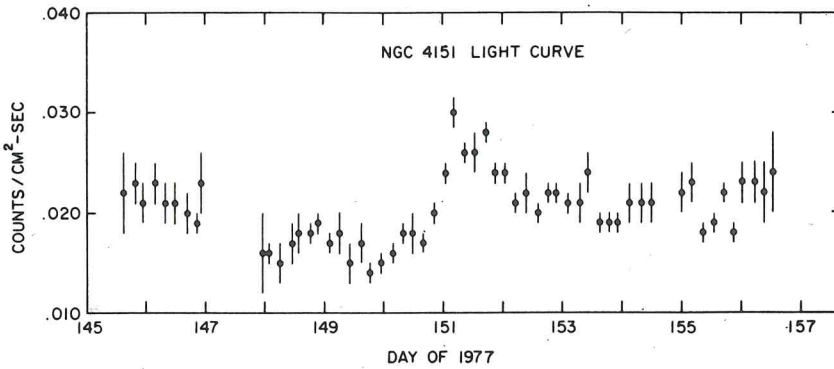


Fig. 34. Light curve of NGC 4151 from OSO-8 data, showing a typical flare-like event (Mushotzky et al, 1978).

sudden or slow transitions in the intensity level. Because of the irregular and unpredictable pattern of the variations in a given source, systematic studies of their incidence on the light curve of objects representative of the various categories could not be pursued so far. However, a fairly general and very important result is that, with the notable exception of the BL Lac's, the various experiments have failed to detect statistically significant spectral variations accompanying the changes of intensity.

The ability of the C/S and the PDS to detect variations in extragalactic objects is illustrated in Fig. 3 and 4 for the case of NGC 4151. The major goals for SAX will be:

i) Detect short time variations during the sightings devoted to the spectral studies. The detection of large values of the rate of change in  $L_x$  can place important constraints on the source structure (radiative transfer effects).

ii) Study systematically the variability on various time scales. These observations are very demanding in time, and will therefore be devoted only to a few representative objects. The WFC will assist in the monitoring of the brightest sources. These studies should be coordinated with regular monitoring programs in other wavebands (from radio to optical). The importance of this coordination is particularly evident for the blazars, where the variations in the soft X-ray component are expected to correlate with those in the optical, while those in the hard component, if dominated by the IC emission, should correlate with the variations at cm-mm wavelengths. The existence of lags or leads between the X-ray variations and those in other bands could also be investigated.

iii) Search for spectral variability in Seyfert's and QSO's. If the lack of evidence so far is due to the smallness of the effect, the very wide range of energies covered by SAX should increase the chances of their detection. Even small spectral changes with the intensity (or their definite absence) are very informative on the emission mechanisms and the degree of homogeneity in the source parameters.

### 3.7. Clusters of Galaxies

Since the discovery by Uhuru that clusters of galaxies are powerful and extended sources of X-rays, and the confirmation by Ariel V and OSO-8 of the thermal nature of the emission with the detection of a strong 6.7 keV iron line, a wealth of spectral and structural data have been collected, particularly with the OSO-8, HEAO-1 and EO instruments. The EO images show beyond doubt that the X-rays come mainly from a diffuse intracluster gas, the HEAO-1 spectra that its iron abundance is typically half-solar, hence that a large fraction of it has been reprocessed through the cluster galaxies. The X-ray data convey information of fun

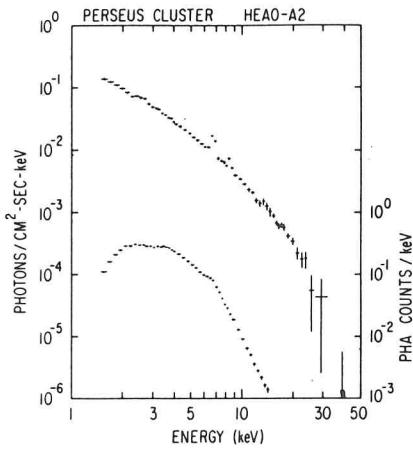


Fig. 35. The spectrum of the Perseus Cluster, showing the  $K\alpha$  and  $K\beta$  iron lines. The lower points are the data before deconvolution (right-hand scale) (Holt and McCray, 1982).

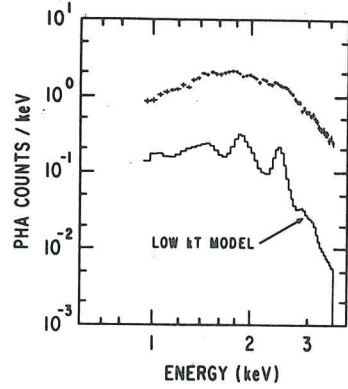


Fig. 36. The SSS spectrum of the core of the Perseus Cluster. Solid line represents best fitting low-temperature model contribution (Mushotzky et al, 1981).

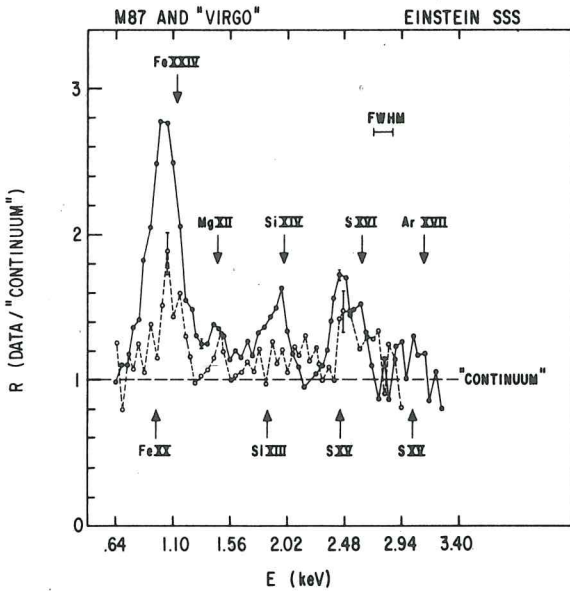


Fig. 37. SSS line emission spectrum of M 87. The dotted line refers to a position  $7'$  away from the center of the galaxy (Lea et al, 1982).

damental importance not only on the distribution, physical state, chemical composition, hence on the history of the intracluster gas, but also on the gravitational potential of all the matter in the clusters, which can be traced through the X-ray maps, hence on their dynamical evolutionary state, and ultimately on the large scale distribution of matter in the Universe.

More than 100 Abell rich clusters, with  $z \leq 0.29$ , have been detected with the HEAO-1 (Johnson et al, 1983), and their  $L_x$  ranges from  $8 \times 10^{42}$  to  $3 \times 10^{45}$  erg/s. Poor clusters with centrally dominant galaxies have  $L_x$  between  $5 \times 10^{41}$  and  $2 \times 10^{44}$  erg/s (EO data, Kriss et al, 1983).

The statistically more accurate medium X-ray energy spectra of the integrated emission from individual rich clusters (Fig.35) can be fitted by a single temperature bremsstrahlung, with  $kT$  typically between 2 and 10 keV, but often two temperature fits are more acceptable, thus indicating the existence of temperature gradients (Mushotzky, 1979). A low  $kT$  component, more pronounced in the more centrally condensed clusters, have been discovered also through the detection of emission lines of O, Si, S below 4 keV with the FPCS and SSS on the EO (Fig.36,37; Canizares et al, 1979; Lea et al, 1982; Mushotzky et al, 1981). This component is interpreted as evidence of a cooling accretion flow in clusters with very deep central potential wells (e.g. Stewart et al, 1983).

The abundance of the heavy elements, the iron in particular, is derived under the assumption that isothermal equilibrium ionization models apply. The iron  $K\alpha/K\beta$ , estimated in a few cases with a (model dependent) spectral inversion of PC data (e.g. Mitchell and Mushotzky, 1980), was found consistent with the above assumption and the temperature derived from the fit of the continuum.

All the detailed spectral studies done so far, for the lack of an adequate spatial resolution, could not investigate closely the temperature and the possible abundance gradients. On the other hand, a detailed broad band surface brightness distribution of many clusters has been obtained with the imaging instruments on the EO. The maps of 46 nearby clusters (Jones and Forman, 1984) have been fitted with an isothermal hydrostatic model (Fig.38) to yield three basic quantities, the core radius  $a$ , the ratio of the energy per unit mass in galaxies and gas, as defined by Cavaliere and Fusco Fermiano (1976),  $\beta = \mu_{\text{H}} \sigma_V^2 / 3kT$ , the magnitude of a central luminosity excess, which is likely to be associated with the low temperature component mentioned above. These quantities, together with  $L_x$ ,  $T$ , the prominence of brightness enhancements around individual galaxies, correlate with the optical morphological properties and are used to produce a classification scheme, with four major categories, clusters of early or evolved type, either with small ( $< 300$  kpc) or large (400-800 kpc) core radii.

Notably the  $\beta$  from the fit is generally less than one, namely the energy per unit mass is greater in the gas than in the galaxies, and various mechanisms to explain this situation have been proposed, including heat conduction from a very hot intergalactic medium (Forman et al, 1984). However in some of the cases where  $\beta$  could be determined independently from the measured galaxy velocity dispersion  $\sigma_v$  and  $kT$ , the two values differ by up to a factor of 2, although the difference is not statistically significant, the quoted errors being very large. This situation may conceal a real discrepancy, due to large scale temperature gradients. The ignorance of these gradients introduces a further uncertainty on the estimates of the gas density and mass. An uncertainty of similar nature affects also the determination of the gravitational mass of M 87 in the Virgo Cluster (Fabricant et al, 1980) and of the dominant galaxies in poor clusters (Kriss et al, 1983).

Changes in the temperature, spatial distribution, chemical composition and X-ray luminosity of the intracluster gas are expected to occur with the cosmic time, as a consequence of the evolution of the cluster potential, the input of metal enriched gas from the galaxies, the heating through dynamical friction, violent cluster relaxation and conduction across the boundary with the intergalactic medium. The EO survey of a considerable sample of optically selected clusters with  $0.1 < z < 0.99$  (Henry et al, 1982), as well as the detailed study of a few high redshift clusters (e.g. White et al, 1981), have failed to find evidence of the predicted evolution, but this failure may result in part from a selection effect (e.g. Forman and Jones, 1982). To avoid the optical selection bias, this type of investigation should be carried out on X-ray selected samples. Furthermore, no reliable information on the temperature of the very distant clusters could be obtained.

From the above it is apparent that the next step in cluster work requires an appropriate combination of spatial resolution, sensitivity and spectral capability. The C/S on SAX is an adequate instrument for the purpose, and some major goals will be pursued, now briefly described.

i) Large scale spectral mapping of nearby clusters. The 46 clusters with  $z < 0.08$  studied by Jones and Forman (1984) have projected core radii ranging from about 1 to 15 arcmin, hence the 2 arcmin spatial resolution of the C/S is adequate for an investigation of temperature and abundance gradients in all of them. The total exposure time for this purpose was estimated around  $10^6$  s by means of simulations. Two of them are given in Fig. 39 and show the raw counts collected in  $1-2 \times 10^4$  s from within the core radius and an annulus 2.5-3.5 core radii of a cluster like Coma ( $kT=8.6$  keV, isothermal model, atomic data from Mewe and Gronenschild, 1981) placed at the maximum distance in the sample,  $z = 0.08$ . In both simulations the temperature and the iron abundance could be estimated with an accuracy better than 15%. This figure illu

strates the ability to resolve the  $K\alpha$  and  $K\beta$  iron lines.

ii) Spectral mapping of low temperature components around dominant galaxies. The exposures just mentioned should also be sufficient to detect low temperature spectral features, hence to evaluate the relevance of cooling flows and their relationship with the magnitude of the central luminosity excess. An additional  $10^6$  s exposure time will be required for a detailed investigation of the most significant cases.

M 87, thanks to its proximity, could be studied in great detail. Fig. 40 illustrates two simulations of high statistical quality, representing the counts collected in  $10^4$  s from a circle  $2'$  in radius and an annulus  $12'-14'$  centered on the galaxy. In poor clusters the X-ray emission is primarily associated with the dominant galaxy, like in the Virgo Cluster, but the distances are substantially larger. The six extended sources of this type analyzed by Kriss et al (1983) lie at distances 5-13 times greater than M 87, and an adequate spectral study would require on average  $10^5$  s for each of them.

iii) Spectral study of high redshift clusters. The confusion limit  $F_c$  of the C/S corresponds to  $z \sim 1$  for a cluster like Coma, hence the spectra of very distant clusters will be measured. The iron line emission can be reliably determined only out to  $z \sim 0.2$ , but the temperature can be measured out to  $z \sim 1$ . To avoid the optical selection bias, a representative sample of clusters distributed between  $0.1 < z < 1$  could be extracted from the All Sky Survey of ROSAT, whose limiting flux is about  $3F_c$ . This selection however will require an extensive deep optical work to secure the identifications of the sources in the ROSAT survey and the redshift measurements, a task which, at least for the faintest objects, could be undertaken with the Space Telescope. A minimum exposure time of  $3 \times 10^6$  s seems required for a sample of about 100 sources.

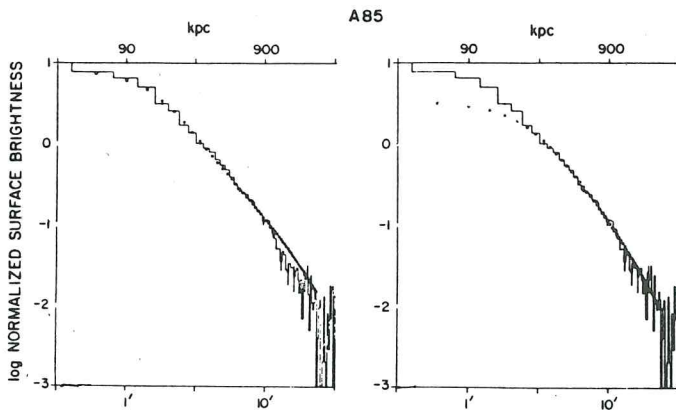


Fig. 38. Surface brightness profile of A85, showing (right) the central luminosity excess above best fitting isothermal model (points) (Jones and Forman, 1984).

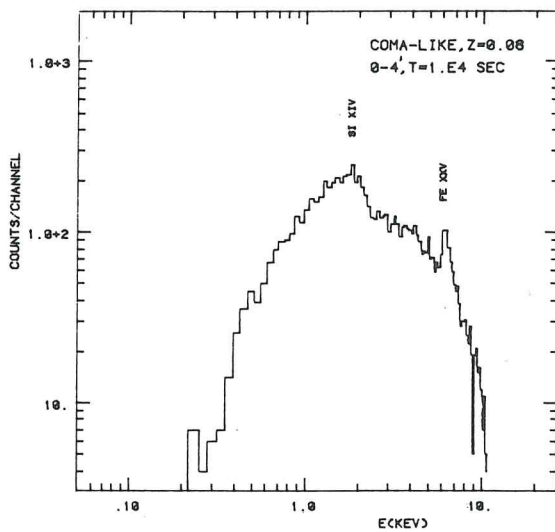


Fig. 39a. Simulation of raw counts collected in  $10^4$  sec with the C/S from Coma-like cluster placed at  $z = 0.08$ , within its core radius. The interstellar absorbing column was taken to be  $10^{21}$  H/cm<sup>2</sup>.

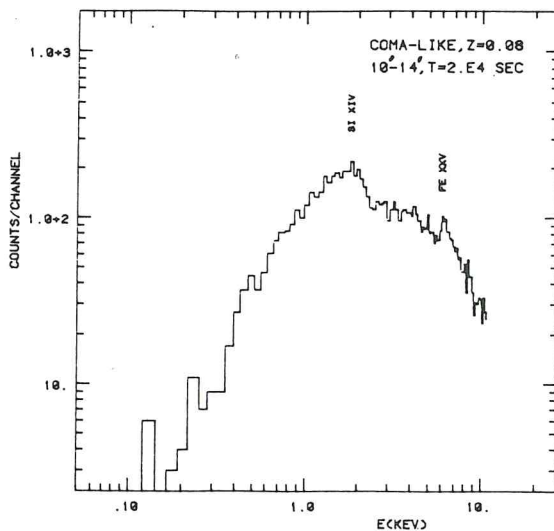


Fig. 39b. Same as Fig. 39a, but counts are from an annulus 2.5-3.5 core radii and exposure time  $2 \times 10^4$  sec.

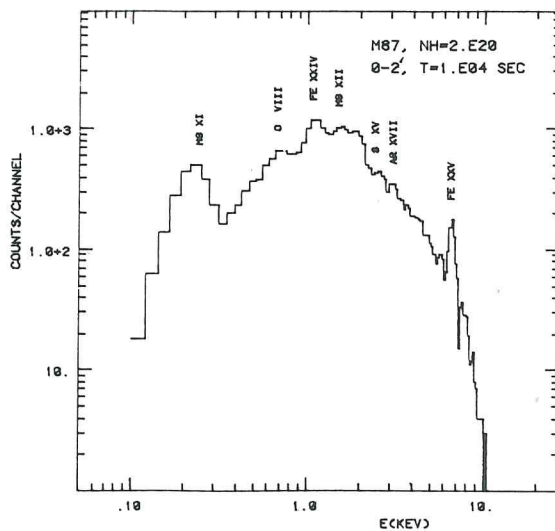


Fig. 40a. Simulation of raw counts collected in  $10^4$  sec with the C/S from a circle  $2'$  radius centered on M 87.

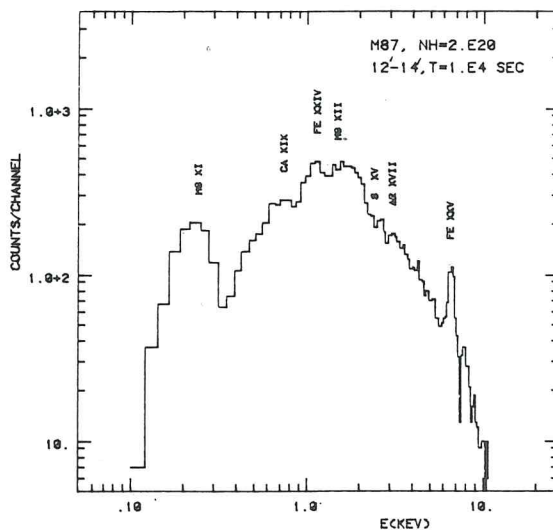


Fig. 40b. Same as Fig. 40a, but counts are from an annulus  $12'$ - $14'$  radius. Temperature gradient adopted from Frabiccant et al (1980).



#### 4. SUMMARY OF OBJECTIVES

The principal objectives are now summarized in Table 2, which contains a sketch of observing program, with the approximate number of sources in the various categories which could be studied in a 2 years mission.

Table 2.

Type of source	Principal objectives	Instruments
<u>X-ray binaries:</u>		
A. BH candidates ( <u>&gt;</u> 6)	Broad band spectroscopy and variability	C/S, PDS, GSPC, (WFC)
B. X-ray pulsars ( <u>&gt;</u> 20)	Pulse and orbital phase study of spectral param. Cyclotron lines	C/S, PDS, GSPC, WFC GSPC, PDS
C. Other binaries (50 - 100)	Broad band spectroscopy and variability. Monitoring of GC sources	C/S, PDS, GSPC, WFC WFC, (C/S)
<u>X-ray transients:</u>		
A. Low latitude	Detection and light curve of bright new and recurrent.	WFC
B. High latitude (100)	Detection, localization, spectral & temp. studies	WFC
<u>Supernova Remnants:</u>		
A. Historical and large SNR (10)	Broad band spectral mapping	C/S, (GSPC)
B. Other galactic and MC remnants (50)	Broad band spectroscopy	C/S
<u>Stellar coronae</u> ( <u>&gt;</u> 100)	Broad band spectroscopy	C/S
<u>Normal galaxies:</u>		
A. Discrete sources in LG gal.s ( <u>&gt;</u> 30)	Broad band spectroscopy	C/S
B. Integrated emiss. of nearby gal. ( <u>&gt;</u> 30)	Broad band spectroscopy	C/S
<u>Active Galactic Nuclei</u>		
A. Bright Seyfert's, QSO's and BL Lac's (50)	Broad band spectroscopy and variability. Long term monitoring of selected obj.	C/S, GSPC, PDS, WFC
B. Faint and very distant AGN (500)	Broad band spectroscopy	C/S

Table 2 (cont.).

Clusters of Galaxies

A. Nearby clusters and groups (50)	Broad band spectral mapping (and HE spectra)	C/S, (GSPC, PDS)
B. Distant clusters (200)	Broad band spectroscopy	C/S

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